

Testing magnetic helicity conservation in a solar-like active event

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ABSTRACT

Context. Magnetic helicity has the remarkable property of being a conserved quantity of ideal magnetohydrodynamics (MHD). Therefore, it could be used as an effective tracer of the magnetic field evolution of magnetised plasmas.

Aims. Theoretical estimations indicate that magnetic helicity is also essentially conserved with non-ideal MHD processes, e.g. magnetic reconnection. This conjecture has however been barely tested, either experimentally or numerically. Thanks to recent advances in magnetic helicity estimation methods, it is now possible to test numerically its dissipation level in general three-dimensional datasets.

Methods. We first revisit the general formulation of the temporal variation of relative magnetic helicity on a fully bounded volume when no hypothesis on the gauge is made. We introduce a method to precisely estimate its dissipation independently of the type of non-ideal MHD processes occurring. In a solar-like eruptive event simulation, using different gauges, we compare its estimation in a finite volume with its time-integrated flux through the boundaries, hence testing the conservation and dissipation of helicity.

Results. We provide an upper bound of the real dissipation of magnetic helicity: It is quasi-null during the quasi-ideal MHD phase. Even when magnetic reconnection is acting the relative dissipation of magnetic helicity is also very small ($< 2.2\%$), in particular compared to the relative dissipation of magnetic energy (> 30 times larger). We finally illustrate how the helicity-flux terms involving velocity components are gauge dependent, hence limiting their physical meaning.

Conclusions. Our study paves the way for more extended and diverse tests of the magnetic helicity conservation properties. Our study confirms the central role that helicity can play in the study of MHD plasmas. For instance, the conservation of helicity can be used to track the evolution of solar magnetic fields, from its formation in the solar interior until their detection as magnetic cloud in the interplanetary space.

Key words. Magnetic fields, Methods: numerical, Sun: surface magnetism, Sun: corona

1. Introduction

In physics, conservation principle have driven the understanding of observed phenomena. Exact and even approximately conserved quantities have allowed to better describe and predict the behaviour of physical systems. Conservation laws state that, for an isolated system, a particular measurable scalar quantity does not change as the system evolves. A corollary is that for a non isolated system, a conserved scalar quantity only evolves thanks to the flux of that quantity through the studied system boundaries. Given a physical paradigm, a physical quantity may not be conserved if source or dissipation terms exist.

In the magnetohydrodynamics (MHD) framework, a quantity has received increasing attention for its conservation property: magnetic helicity (Elsasser 1956). Magnetic helicity quantitatively describes the geometrical degree of twist, shear, or more generally, knottedness of magnetic field lines (Moffatt 1969). In ideal MHD, where magnetic field can be described as the collection of individual magnetic field lines, magnetic helicity is a strictly conserved quantity (Woltjer 1958) as no dissipation, nor creation, of helicity is permitted since magnetic field line cannot reconnect.

In his seminal work, Taylor (1974) conjectured that even in non-ideal MHD, the dissipation of magnetic helicity should be relatively weak, and hence, that magnetic helicity should be conserved. This could be theoretically explained by the inverse-cascade property of magnetic helicity: in turbulent medium, helicity unlike magnetic energy, tends to cascade towards the larger spatial scales, thus avoiding dissipation at smaller scales (Frisch et al. 1975; Pouquet et al. 1976). This cascade has been observed in numerical simulations (Alexakis et al. 2006; Mininni 2007) as well as in laboratory experiments (Ji et al. 1995). In the case of resistive MHD, Berger (1984) derived an upper limit on the amount of magnetic helicity that could be dissipated through constant resistivity. He showed that the typical helicity dissipation time in the solar corona was far exceeding the one for magnetic energy dissipation.

From the Taylor's conjecture on helicity conservation have been derived multiple important consequences for the dynamics of plasma systems. Based on helicity conservation, Taylor (1974) predicted that relaxing MHD systems should reach a linear force free state. This prediction, which was verified to different degrees, has allowed to understand the dynamics of plasma in several laboratory experiments (Taylor

1986; Prager 1999; Ji 1999; Yamada 1999). The concept of Taylor relaxation has also driven theoretical models of solar coronal heating (Heyvaerts & Priest 1984). The importance of magnetic helicity conservation has been further raised for magnetic field dynamos (Brandenburg & Subramanian 2005; Blackman & Hubbard 2014). Magnetic helicity is also strongly impacting the energy budget during reconnection events (Linton et al. 2001; Linton & Antiochos 2002; Del Sordo et al. 2010), and models of eruption based on magnetic helicity annihilation have been developed (Kusano et al. 2004). The conservation of magnetic helicity has been suggested as the core reason of the existence of Coronal Mass Ejections (CMEs), the latter being the mean for the Sun to expel its excess magnetic helicity (Rust 1994; Low 1996). Because of this hypothesis, important efforts to estimate the magnetic helicity in the solar coronal have been carried over the last decades (Démoulin 2007; Démoulin & Pariat 2009).

Despite its potential importance, it is surprising to note that Taylor's conjecture on helicity conservation has been barely tested. Several experimental and numerical works have focused on testing Taylor's prediction that relaxed systems should be linear force free (Taylor 1986; Prager 1999; Ji 1999; Yamada 1999; Antiochos & DeVore 1999). However, the fact that the system may not reach a complete relaxed linear force free state, does not mean that helicity is not well conserved. It simply means that the dynamics of the relaxation does not allow the full redistribution of helicity towards the largest available scales. Looking at the long term evolution of turbulent MHD systems, magnetic helicity has been found to decay slower than magnetic energy (Biskamp & Müller 1999; Christensson et al. 2005; Candelaresi 2012). These studies are however based on the estimation of the helicity density which is not generally gauge invariant. The periodic boundaries, the long term and large spatial scales involved do not allow an estimation of helicity dissipation when a "singular" non-ideal event is occurring.

Direct tests on magnetic helicity conservation and dissipation have actually been limited because of the inherent difficulties to measure that very quantity. Laboratory experiments and observations requires the measurement of the full 3D distribution of the magnetic field. Most laboratory experiments therefore make assumptions on the symmetries of the system to limit the sampling of the data, hence limiting measurement precision. Ji et al. (1995) have measured a 1.3–5.1% decay of helicity (relatively to a 4–10% decrease of energy) during a sawtooth relaxation in a reversed-field pinch experiment. Heidbrink & Dang (2000) have estimated that helicity was conserved within 1% during a sawtooth crash. Other experiments have tried to measure helicity conservation by comparing the helicity in the system with its theoretically injected amount (Barnes et al. 1986; Gray et al. 2010). While they found results agreeing within 10–20%, the experimental conditions limit the precision on the measure of injected helicity (Stallard et al. 2003). In numerical experiments of flux tubes reconnection, using triple periodic boundaries, Linton & Antiochos (2005) have found that the loss of helicity was ranging between 6% and 53% depending on the Lundquist number, hence primarily due to diffusion rather than magnetic reconnection.

For non-periodic systems where the magnetic field is threading the domain boundary (as in most natural cases), the very definition of magnetic helicity is not gauge invariant (cf. Sect. 2.1), and a modified definition of magnetic helicity, the relative magnetic helicity, had to be introduced (Berger & Field 1984, see Sect. 2.2). Only very recently, practical methods that allow to generally computed relative magnetic helicity have

been published (Rudenko & Myshyakov 2011; Thalmann et al. 2011; Valori et al. 2012; Yang et al. 2013). These methods are now opening the possibilities to carefully study helicity of 3D datasets such as the ones frequently used for natural plasmas. Zhao et al. (2015) and Knizhnik et al. (2015) have observed a good helicity conservation but have not quantified it. Yang et al. (2013) have measured the helicity conservation in a numerical simulation by comparing the relative helicity flux with the variation of helicity in the domain. They found that during the quasi-ideal evolution of the system the helicity was conserved within 3%, with higher dissipation values when non-ideal MHD effects become important. However, in previous studies, the relative helicity flux definition used may not be fully consistent with the choice of gauge used to compute the volume helicity (see Sect. 2.2).

In the present manuscript, we will further push the test on the conservation of magnetic helicity. We will derive a generalized analytical formula for the flux of relative magnetic helicity (Sect. 2.3) without taking any assumption on the gauges of the studied and the reference fields. We will discuss whether relative magnetic helicity can be considered as a conserved quantity in the classical sense, that is, whether its variations can be described solely as a flux through the boundary. The general method that we will employ (see Sect. 3) is based on the comparison of the evolution of the relative magnetic helicity with its flux through the boundaries, and will be applied to a numerical simulation of solar active-like events. In Sections 4 & 5, using different gauges we will constraint the level of conservation of relative magnetic helicity, and study the amount of dissipated magnetic helicity. This is further extended in Appendix A with another selection of gauges. We will finally conclude in Sect. 6

2. Magnetic helicity and its time variation

2.1. Magnetic helicity

Magnetic helicity is defined as

$$H_m = \int_{\mathcal{V}} \mathbf{A} \cdot \mathbf{B} d\mathcal{V}, \quad (1)$$

where \mathbf{B} is the magnetic field studied over a fixed volume \mathcal{V} . \mathcal{V} is here a fully closed¹ volume bounded by the surface $\partial\mathcal{V}$. The vector potential, \mathbf{A} of \mathbf{B} classically verifies: $\nabla \times \mathbf{A} = \mathbf{B}$ as $\nabla \cdot \mathbf{B} = 0$ (only approximately true in numerical computations, Valori et al. 2013). The magnetic field \mathbf{B} is gauge invariant, meaning that it is unchanged by transformations $\mathbf{A} \rightarrow \mathbf{A} + \nabla\psi$, where ψ is any sufficiently regular, scalar function of space and time. Since \mathbf{A} is not uniquely defined in general, magnetic helicity requires additional constraints in order to be well-defined. In particular, H_m is a gauge-invariant quantity provided that \mathcal{V} is a magnetic volume, i.e. that the magnetic field is tangent at any point of the surface boundary $\partial\mathcal{V}$ of \mathcal{V} : $(\mathbf{B} \cdot d\mathbf{S})|_{\partial\mathcal{V}} = 0$, at any time.

Assuming that the volume \mathcal{V} is fixed, with $\partial\mathcal{V}$ a flux surface ensuring gauge invariance, the temporal variation of magnetic helicity is derived by direct differentiation of Eq. (1):

$$\frac{dH_m}{dt} = \int_{\mathcal{V}} \frac{\partial \mathbf{A}}{\partial t} \cdot \mathbf{B} d\mathcal{V} + \int_{\mathcal{V}} \mathbf{A} \cdot \frac{\partial \mathbf{B}}{\partial t} d\mathcal{V}, \quad (2)$$

¹ We do not consider the particular cases of infinite volumes which require specific hypothesis on the behaviour of the quantities at large distance.

where each of the two integrals on the right hand side are well defined, since they are independently gauge invariant. Given that $\nabla \cdot (\mathbf{A} \times \partial \mathbf{A} / \partial t) = \partial \mathbf{A} / \partial t \cdot \nabla \times \mathbf{A} - \mathbf{A} \cdot \partial (\nabla \times \mathbf{A}) / \partial t$ and using the Gauss divergence theorem, one obtains:

$$\frac{dH_m}{dt} = \int_{\partial \mathcal{V}} \left(\mathbf{A} \times \frac{\partial \mathbf{A}}{\partial t} \right) \cdot d\mathbf{S} + 2 \int_{\mathcal{V}} \mathbf{A} \cdot \frac{\partial \mathbf{B}}{\partial t} d\mathcal{V}.$$

Here, $d\mathbf{S}$ is the elementary surface vector, directed outside of the domain \mathcal{V} . Using the Faraday's law of induction, one derives:

$$\frac{dH_m}{dt} = \int_{\partial \mathcal{V}} \left(\mathbf{A} \times \frac{\partial \mathbf{A}}{\partial t} \right) \cdot d\mathbf{S} - 2 \int_{\mathcal{V}} \mathbf{A} \cdot \nabla \times \mathbf{E} d\mathcal{V}, \quad (3)$$

Using again the Gauss divergence theorem, one finds that the temporal variation of magnetic helicity is composed of three independently gauge-invariant terms, *i.e.*, a volume dissipative term and two helicity flux terms on the surface of \mathcal{V} , such that:

$$\frac{dH_m}{dt} = \left. \frac{dH_m}{dt} \right|_{\text{diss}} + F_{m,B} + F_{m,A} \quad \text{with} \quad (4)$$

$$\left. \frac{dH_m}{dt} \right|_{\text{diss}} = -2 \int_{\mathcal{V}} \mathbf{E} \cdot \mathbf{B} d\mathcal{V} \quad (5)$$

$$F_{m,B} = 2 \int_{\partial \mathcal{V}} (\mathbf{A} \times \mathbf{E}) \cdot d\mathbf{S} \quad (6)$$

$$F_{m,A} = \int_{\partial \mathcal{V}} \left(\mathbf{A} \times \frac{\partial \mathbf{A}}{\partial t} \right) \cdot d\mathbf{S} \quad (7)$$

In ideal MHD, where $\mathbf{E} = -\mathbf{v} \times \mathbf{B}$, the volume term is null. For an isolated system, magnetic helicity is thus conserved in the classical sense since its variations are null (cf. Sect. 1). Variations of H_m can only originate from advection of helicity through the boundaries of \mathcal{V} . The $dH_m/dt|_{\text{diss}}$ term corresponds to the dissipation of magnetic helicity of the studied magnetic field in \mathcal{V} . Taylor (1974) conjectured that this term is relatively small even when non-ideal MHD processes are developing, *e.g.*, when magnetic reconnection is present (cf. Sect. 1).

However, because of the gauge invariance requirement which imposes that $\partial \mathcal{V}$ must be a flux surface, magnetic helicity appears as a quantity of limited practical use. In most studied systems, magnetic field is threading the surface $\partial \mathcal{V}$ and the condition $(\mathbf{B} \cdot d\mathbf{S})|_{\partial \mathcal{V}} = 0$ is not fulfilled. In their seminal paper, Berger & Field (1984) have introduced the concept of relative magnetic helicity: a gauge-invariant quantity which preserve essential properties of magnetic helicity while allowing a non null normal component of the field \mathbf{B} through the surface of the studied domain.

2.2. Relative Magnetic helicity

In their initial work, Berger & Field (1984) gave a first definition of the relative magnetic helicity as the difference between the helicity of the studied field \mathbf{B} and the helicity of a reference field \mathbf{B}_0 having the same distribution than \mathbf{B} of the normal component along the surface: $((\mathbf{B}_0 - \mathbf{B}) \cdot d\mathbf{S})|_{\partial \mathcal{V}} = 0$.

While the definition allows for any field to be used as the reference field, the potential field \mathbf{B}_p is frequently used as a reference field in the literature. Since $\nabla \times \mathbf{B}_p = 0$, the potential field can be derived from a scalar function ϕ :

$$\mathbf{B}_p = \nabla \phi \quad (8)$$

where the scalar potential ϕ is the solution of the Laplace equation $\Delta \phi = 0$ derived from $\nabla \cdot \mathbf{B}_p = 0$. Given the distribution of

the normal component on the surface $\mathbf{B} \cdot d\mathbf{S}|_{\partial \mathcal{V}} = \partial \phi / \partial n$ of the studied domain, at any instant there is a unique potential field which satisfies the following condition on the whole boundary of the volume considered:

$$(\mathbf{B}_p \cdot d\mathbf{S})|_{\partial \mathcal{V}} = (\mathbf{B} \cdot d\mathbf{S})|_{\partial \mathcal{V}} \quad (9)$$

Under the above assumptions, the potential field has the lowest possible energy for the given distribution of \mathbf{B} on $\partial \mathcal{V}$ (*e.g.* Eq.(2) of Valori et al. 2012). In the following we will also use the potential field as our reference field.

A second gauge-independent definition for relative magnetic helicity has been given by Finn & Antonsen (1985), definition used from here on in this article:

$$H = \int_{\mathcal{V}} (\mathbf{A} + \mathbf{A}_p) \cdot (\mathbf{B} - \mathbf{B}_p) d\mathcal{V}. \quad (10)$$

with \mathbf{A}_p the potential vector of the potential field $\mathbf{B}_p = \nabla \times \mathbf{A}_p$. Not only is H gauge invariant, but the gauges of \mathbf{A} and \mathbf{A}_p are independent of each others, *i.e.*, for any set of sufficiently-regular scalar functions $(\psi; \psi_p)$, H will be unchanged by the gauges transformation $(\mathbf{A}; \mathbf{A}_p) \rightarrow (\mathbf{A} + \nabla \psi; \mathbf{A}_p + \nabla \psi_p)$. Let us note that \mathbf{A}_p and ϕ correspond to two distinct solutions of the Helmholtz's theorem, *i.e.* two distinct non-incompatible decompositions of \mathbf{B}_p .

The relative helicity in Eq. (10) can first be decomposed in a contribution due only to \mathbf{B} , Eq. (1), one only to \mathbf{B}_p , and a mixed term:

$$H = H_m - H_p + H_{\text{mix}} \quad \text{with} \quad (11)$$

$$H_p = \int_{\mathcal{V}} \mathbf{A}_p \cdot \mathbf{B}_p d\mathcal{V} \quad (12)$$

$$H_{\text{mix}} = \int_{\mathcal{V}} (\mathbf{A}_p \cdot \mathbf{B} - \mathbf{A} \cdot \mathbf{B}_p) d\mathcal{V} = \int_{\partial \mathcal{V}} (\mathbf{A} \times \mathbf{A}_p) \cdot d\mathbf{S} \quad (13)$$

Let us note that this decomposition is only formal in the sense that each term is actually gauge dependent and only their sum is gauge invariant.

Relative magnetic helicity, as defined in Eq. (10), is equal to the difference between the helicity of the field \mathbf{B} and the helicity of its potential field ($H = H_m(\mathbf{B}) - H_m(\mathbf{B}_p)$) only if H_{mix} cancels. Relative helicity is in general not a simple difference of helicity, as for relative energy. A sufficient (but not necessary) condition that ensures the nullity of the mixed term is that \mathbf{A} and \mathbf{A}_p have the same transverse component on the surface, *i.e.*:

$$\mathbf{A} \times d\mathbf{S}|_{\partial \mathcal{V}} = \mathbf{A}_p \times d\mathbf{S}|_{\partial \mathcal{V}} \quad (14)$$

This condition automatically enforces the condition of Eq. (9) on the normal field components. However it imposes to link the choice of the gauge of \mathbf{A} with that of \mathbf{A}_p . The original definition of Berger & Field (1984) corresponds to a quantity which is less general than the one given by Finn & Antonsen (1985). It is only gauge invariant for particular sets of transformation: $(\mathbf{A}; \mathbf{A}_p) \rightarrow (\mathbf{A} + \nabla \psi; \mathbf{A}_p + \nabla \psi_p)$.

Another possible decomposition of relative magnetic helicity from Eq. (10) is (Berger 2003):

$$H = H_j + 2H_{\text{pj}} \quad \text{with} \quad (15)$$

$$H_j = \int_{\mathcal{V}} (\mathbf{A} - \mathbf{A}_p) \cdot (\mathbf{B} - \mathbf{B}_p) d\mathcal{V} \quad (16)$$

$$H_{\text{pj}} = \int_{\mathcal{V}} \mathbf{A}_p \cdot (\mathbf{B} - \mathbf{B}_p) d\mathcal{V} \quad (17)$$

where H_j is the classical magnetic helicity of the non-potential, or current carrying, component of the magnetic field, $\mathbf{B}_j = \mathbf{B} - \mathbf{B}_p$ and H_{pj} is the mutual helicity between \mathbf{B}_p and \mathbf{B}_j . The field \mathbf{B}_j is contained within the volume \mathcal{V} so it is also called the closed field part of \mathbf{B} . Because of Eq. (9), not only H , but also both H_j and H_{pj} are independently gauge invariant.

2.3. Relative magnetic helicity variation

Assuming a fixed domain \mathcal{V} , we can differentiate Eq. (10) in time in order to study the time variations of relative helicity:

$$\begin{aligned} \frac{dH}{dt} &= \int_{\mathcal{V}} \frac{\partial(\mathbf{A} + \mathbf{A}_p)}{\partial t} \cdot \nabla \times (\mathbf{A} - \mathbf{A}_p) d\mathcal{V} \\ &+ \int_{\mathcal{V}} (\mathbf{A} + \mathbf{A}_p) \cdot \frac{\partial(\mathbf{B} - \mathbf{B}_p)}{\partial t} d\mathcal{V} \end{aligned} \quad (18)$$

Using the Gauss divergence theorem one obtains:

$$\begin{aligned} \frac{dH}{dt} &= \int_{\partial\mathcal{V}} \left((\mathbf{A} - \mathbf{A}_p) \times \frac{\partial(\mathbf{A} + \mathbf{A}_p)}{\partial t} \right) \cdot d\mathbf{S} \\ &+ \int_{\mathcal{V}} (\mathbf{A} - \mathbf{A}_p) \cdot \frac{\partial(\mathbf{B} + \mathbf{B}_p)}{\partial t} d\mathcal{V} \\ &+ \int_{\mathcal{V}} (\mathbf{A} + \mathbf{A}_p) \cdot \frac{\partial(\mathbf{B} - \mathbf{B}_p)}{\partial t} d\mathcal{V} \end{aligned}$$

Combining the second and third terms, one finds the following synthetic decomposition of the helicity variation in three terms:

$$\begin{aligned} \frac{dH}{dt} &= 2 \int_{\mathcal{V}} \mathbf{A} \cdot \frac{\partial \mathbf{B}}{\partial t} d\mathcal{V} + \int_{\partial\mathcal{V}} \left((\mathbf{A} - \mathbf{A}_p) \times \frac{\partial(\mathbf{A} + \mathbf{A}_p)}{\partial t} \right) \cdot d\mathbf{S} \\ &- 2 \int_{\mathcal{V}} \mathbf{A}_p \cdot \frac{\partial \mathbf{B}_p}{\partial t} d\mathcal{V} \end{aligned} \quad (19)$$

This decomposition is only formal. Indeed, as for the decomposition of relative helicity of Eq. (11), none of these three terms are independently gauge invariant and only their sum is. The third term can be further decomposed using the scalar potential ϕ , Eq. (8), and the Gauss divergence theorem:

$$\begin{aligned} \frac{dH}{dt} \Big|_{B_p} &= -2 \int_{\mathcal{V}} \mathbf{A}_p \cdot \frac{\partial \mathbf{B}_p}{\partial t} d\mathcal{V} = -2 \int_{\mathcal{V}} \mathbf{A}_p \cdot \nabla \left(\frac{\partial \phi}{\partial t} \right) d\mathcal{V} \\ &= -2 \int_{\partial\mathcal{V}} \frac{\partial \phi}{\partial t} \mathbf{A}_p \cdot d\mathbf{S} + 2 \int_{\mathcal{V}} \frac{\partial \phi}{\partial t} \nabla \cdot \mathbf{A}_p d\mathcal{V} \end{aligned} \quad (20)$$

Using the Faraday law and the Gauss divergence theorem we also obtain:

$$\int_{\mathcal{V}} \mathbf{A} \cdot \frac{\partial \mathbf{B}}{\partial t} d\mathcal{V} = - \int_{\partial\mathcal{V}} (\mathbf{E} \times \mathbf{A}) \cdot d\mathbf{S} - \int_{\mathcal{V}} \mathbf{B} \cdot \mathbf{E} d\mathcal{V} \quad (21)$$

Assuming that at the boundary the evolution of the system is ideal: $\mathbf{E}|_{\partial\mathcal{V}} = (-\mathbf{v} \times \mathbf{B})|_{\partial\mathcal{V}}$ the surface flux can be written as (e.g. Berger & Field 1984):

$$- \int_{\partial\mathcal{V}} (\mathbf{E} \times \mathbf{A}) \cdot d\mathbf{S} = - \int_{\partial\mathcal{V}} (\mathbf{B} \cdot \mathbf{A}) \mathbf{v} \cdot d\mathbf{S} + \int_{\partial\mathcal{V}} (\mathbf{v} \cdot \mathbf{A}) \mathbf{B} \cdot d\mathbf{S} \quad (22)$$

Note that if the evolution of the system was not ideal at the boundary, an additional flux term depending on $\mathbf{E}_{\text{non ideal}} \times \mathbf{A}$, could be added, with $\mathbf{E}_{\text{non ideal}}$ the non ideal part of the electric field. In the study presented here this term will de facto be estimated but assumed to be measured as a volume dissipation term.

Including Eqs. (20 - 22) in Eq. (19), the variation of magnetic helicity can thus be decomposed as:

$$\frac{dH}{dt} = \frac{dH}{dt} \Big|_{\text{diss}} + \frac{dH}{dt} \Big|_{B_p, \text{var}} + F_{Vn} + F_{Bn} + F_{AAp} + F_{\phi} \quad (23)$$

with

$$\frac{dH}{dt} \Big|_{\text{diss}} = -2 \int_{\mathcal{V}} \mathbf{E} \cdot \mathbf{B} d\mathcal{V} \quad (24)$$

$$\frac{dH}{dt} \Big|_{B_p, \text{var}} = 2 \int_{\mathcal{V}} \frac{\partial \phi}{\partial t} \nabla \cdot \mathbf{A}_p d\mathcal{V} \quad (25)$$

$$F_{Vn} = -2 \int_{\partial\mathcal{V}} (\mathbf{B} \cdot \mathbf{A}) \mathbf{v} \cdot d\mathbf{S} \quad (26)$$

$$F_{Bn} = 2 \int_{\partial\mathcal{V}} (\mathbf{v} \cdot \mathbf{A}) \mathbf{B} \cdot d\mathbf{S} \quad (27)$$

$$F_{AAp} = \int_{\partial\mathcal{V}} \left((\mathbf{A} - \mathbf{A}_p) \times \frac{\partial(\mathbf{A} + \mathbf{A}_p)}{\partial t} \right) \cdot d\mathbf{S} \quad (28)$$

$$F_{\phi} = -2 \int_{\partial\mathcal{V}} \frac{\partial \phi}{\partial t} \mathbf{A}_p \cdot d\mathbf{S} \quad (29)$$

The $dH/dt|_{\text{diss}}$ term is a volume term which corresponds to the actual dissipation of magnetic helicity of the studied magnetic field (Eq. 5). The $dH/dt|_{B_p, \text{var}}$ term, which, despite being a volume term, is not a dissipation, traces a change in the helicity of the potential field. As \mathbf{B} is evolving, its distribution at the boundary implies a changing \mathbf{B}_p (Eq. 9). The helicity of the potential field, not necessarily null, is therefore evolving in time. More precisely, the potential field is defined only in terms of its boundary values. However, this is not true in general for the vector potential of the potential field, because of the gauge freedom. Hence, in general, the helicity of the potential field cannot be expressed as a function of boundary values only, except for the particular case of a vector potential without sources or sinks in \mathcal{V} , *i.e.*, when the Coulomb gauge is used. Therefore, the time variation of the helicity of the potential field necessarily contains both volume and flux contributions.

All the other terms, are flux terms that correspond to the transfer of helicity through the surface boundary $\partial\mathcal{V}$. The F_{Vn} and F_{Bn} are sometimes called the "emergence" and "shear" terms, but such a characterization can be misleading as their contributions depend on the gauge selected for \mathbf{A} . The F_{AAp} term is related to a cross contribution of \mathbf{A} and \mathbf{A}_p . Finally F_{ϕ} correspond to a flux of the helicity of the potential field.

The $dH/dt|_{\text{diss}}$ term is the only term of the decomposition which is gauge invariant. All the other terms are not independently gauge invariant. It means that the relative intensity of these terms will be different for different gauges. Combined, they produce the same, gauge-invariant value of dH/dt . We will study the dependence of the above decomposition on the chosen gauge in Sect. 5. We should also note that the total flux F_{tot} of relative helicity,

$$F_{\text{tot}} = F_{Vn} + F_{Bn} + F_{AAp} + F_{\phi}, \quad (30)$$

is only gauge invariant for \mathbf{A} but not for \mathbf{A}_p .

Moreover, unlike $dH/dt|_{\text{diss}}$, $dH/dt|_{B_p, \text{var}}$ is not a priori null in ideal MHD. It implies that dH/dt cannot be written in a classical conservative form since $dH/dt|_{B_p, \text{var}}$ cannot be strictly written as a flux term. Therefore relative magnetic helicity is not a priori a conserved quantity of MHD in the classical sense: its variation in \mathcal{V} may not solely come from a flux of relative helicity through the boundary. Relative magnetic helicity may not be conserved

even if $dH/dt|_{\text{diss}}$ is small. The conservation of relative helicity and the dissipation of magnetic helicity are thus two distinct problems: the relative intensity of these terms will vary, depending on which gauge is employed as we will illustrate in Sect. 5.

2.4. Relative Magnetic helicity variation with specific gauge conditions

While the variation of magnetic helicity can be generally described by Eq. (23) for any gauge, the choice of some specific additional constraint on the gauge allows to simplify the expression of dH/dt and possibly its computation.

2.4.1. The $\mathbf{A}|_{\partial\mathcal{V}} = \mathbf{A}_p|_{\partial\mathcal{V}}$ condition

We note that with the specific condition

$$\mathbf{A}|_{\partial\mathcal{V}} = \mathbf{A}_p|_{\partial\mathcal{V}}, \quad (31)$$

the condition of Eq. (14) is necessarily satisfied thus $F_{\text{AAp}} = 0$ and the terms F_{Vn} and F_{Bn} can be expressed only in terms of \mathbf{A}_p . Thus, the helicity variation, Eq. (23), simplifies as:

$$\begin{aligned} \left. \frac{dH}{dt} \right|_{\text{Cond.}(31)} = & -2 \int_{\mathcal{V}} \mathbf{E} \cdot \mathbf{B} d\mathcal{V} + 2 \int_{\mathcal{V}} \frac{\partial\phi}{\partial t} \nabla \cdot \mathbf{A}_p d\mathcal{V} \\ & - 2 \int_{\partial\mathcal{V}} (\mathbf{B} \cdot \mathbf{A}_p) \mathbf{v} \cdot d\mathbf{S} + 2 \int_{\partial\mathcal{V}} (\mathbf{v} \cdot \mathbf{A}_p) \mathbf{B} \cdot d\mathbf{S} \\ & - 2 \int_{\partial\mathcal{V}} \frac{\partial\phi}{\partial t} \mathbf{A}_p \cdot d\mathbf{S} \end{aligned} \quad (32)$$

We note that in this derivation the vector potential \mathbf{A} is absent. The condition (31) allows to get rid of the need of computing \mathbf{A} to estimate the helicity variations. However, to derive H from Eq. (10) both \mathbf{A} and \mathbf{A}_p must be computed with gauges coupled with Eq. (31). Then, one must strictly control that this condition is enforced all over the surface of the studied system. This can actually be numerically challenging.

In the present manuscript, in order to determine the helicity dissipation, we will compare time integrated helicity flux with direct helicity measurements. As our numerical method does not allow us to enforce Eq. (31), we cannot use the simplified Eq. (32) to compute the helicity variation.

2.4.2. Coulomb gauge for potential field: $\nabla \cdot \mathbf{A}_p = 0$

If to determine the potential field one decides to choose the Coulomb gauge,

$$\nabla \cdot \mathbf{A}_p = 0, \quad (33)$$

then there is no volume variation of the helicity of the potential field $dH/dt|_{\text{Bp,var}} = 0$. The helicity variation can be reduced to the simplified form:

$$\left. \frac{dH}{dt} \right|_{\text{Cond.}(33)} = -2 \int_{\mathcal{V}} \mathbf{B} \cdot \mathbf{E} d\mathcal{V} + F_{\text{tot}} \quad (34)$$

Using the coulomb gauge for the potential field one observes that the variation of the relative magnetic helicity are given by a flux of helicity through the boundary and the dissipation term $dH/dt|_{\text{diss}}$. Relative magnetic helicity to a reference field expressed in the Coulomb gauge can therefore be written with a classical conservative equation.

2.4.3. Boundary null Coulomb gauge

The condition (33) does not enforce a unique solution for \mathbf{A}_p . It is possible to further constrain the Coulomb gauge if the vector potential \mathbf{A}_p satisfies the additional boundary condition:

$$\mathbf{A}_p \cdot d\mathbf{S}|_{\partial\mathcal{V}} = 0. \quad (35)$$

With this further constraint the flux of helicity of the potential field is null, $F_\phi = 0$. Then, with conditions (33) and (35), the helicity variation thus reduces to the form:

$$\begin{aligned} \left. \frac{dH}{dt} \right|_{\text{Cond.}(35)} = & -2 \int_{\mathcal{V}} \mathbf{E} \cdot \mathbf{B} d\mathcal{V} + \int_{\partial\mathcal{V}} \left(2(\mathbf{v} \cdot \mathbf{A}) \mathbf{B} \right. \\ & \left. - 2(\mathbf{B} \cdot \mathbf{A}) \mathbf{v} + (\mathbf{A} - \mathbf{A}_p) \times \frac{\partial(\mathbf{A} + \mathbf{A}_p)}{\partial t} \right) \cdot d\mathbf{S} \end{aligned} \quad (36)$$

2.4.4. Simplified helicity flux

Including the three conditions of previous sub-sections, *i.e.*,

$$\begin{cases} \mathbf{A}|_{\partial\mathcal{V}} = \mathbf{A}_p|_{\partial\mathcal{V}} \\ \nabla \cdot \mathbf{A}_p = 0 \\ \mathbf{A}_p \cdot d\mathbf{S}|_{\partial\mathcal{V}} = 0. \end{cases} \quad (37)$$

we obtain the well known expression for the simplified helicity flux (e.g. Berger & Field 1984; Pariat et al. 2005):

$$\begin{aligned} \left. \frac{dH}{dt} \right|_{\text{Cond.}(37)} = & -2 \int_{\mathcal{V}} \mathbf{E} \cdot \mathbf{B} d\mathcal{V} \\ & + 2 \int_{\partial\mathcal{V}} ((\mathbf{v} \cdot \mathbf{A}_p) \mathbf{B} - (\mathbf{B} \cdot \mathbf{A}_p) \mathbf{v}) \cdot d\mathbf{S} \end{aligned} \quad (38)$$

With this set of conditions, the flux terms F_{Vn} and F_{Bn} are fully fixed. However, one should remind that these terms remain gauge dependent: using a gauge where the conditions of Eq. (37) are not fully enforced would lead to a different \mathbf{A}_p and \mathbf{A} , consequently, to a different distribution of helicity flux between F_{Vn} , F_{Bn} and other terms. It is therefore theoretically incorrect to study them independently.

Equation (38) is the classical formulation for the helicity flux that has been derived by Berger & Field (1984) for an infinite plane. However in the case of a 3D cubic domain this formulation is only valid if all the conditions of Eq. (37) are satisfied. While Yang et al. (2013) have used Eq. (38) to compute the helicity flux, it remains to be determined if all conditions (37) are fulfilled when computing the volume helicity. Their relatively high level of non-conservation (3%) in the ideal phase of the evolution of their system may be related to this discrepancy. Finally, while conditions of Eq. (37) can drastically simplify the estimation of the helicity variation, they strongly constraint the numerical implementation of \mathbf{A} and \mathbf{A}_p . Fast, precise and practical numerical computation of the vector potentials may require a different choice a gauge.

3. Methodology

The present section describes the methodology employed in our numerical experiments. In Sect. 3.1 we present the general method used to estimate the relative magnetic helicity conservation and the magnetic helicity dissipation. Then in Sect. 3.2 we describe how we practically compute the volume helicity and its flux. Finally, in Sect. 3.3 we present the numerical data set considered for the helicity conservation tests.

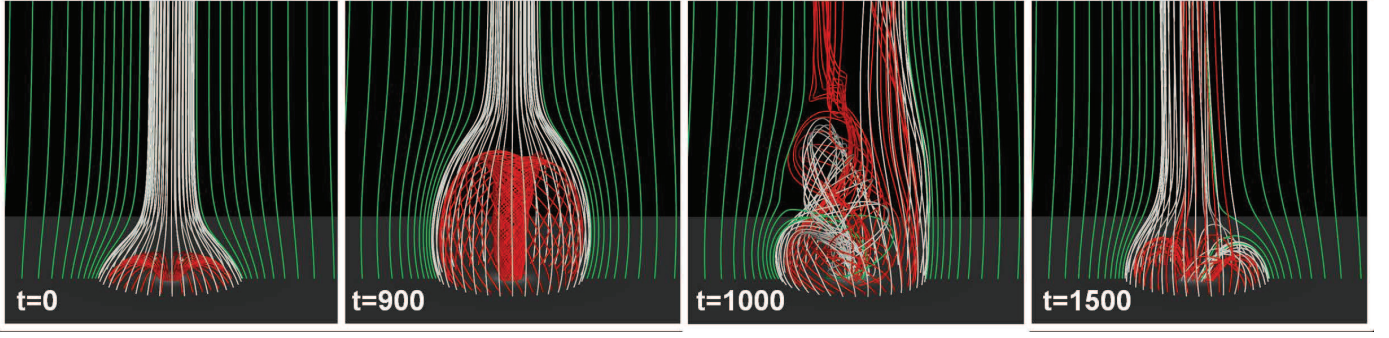


Fig. 1. Snapshots of the magnetic field evolution during the generation of the jet. The red field lines are initially closed. The green and white field lines are initially open. All the field lines are plotted from fixed footpoints. The red and white field lines are regularly plotted along a circle of constant radius while the green field lines are plotted along the x axis. At $t = 900$ the system is in its pre-eruption stage. It is close to the maximum of energy and helicity. All the helicity is stored in the close domain. At $t = 1000$ the system is erupting. Numerous field lines have changed of connectivity as one observe open red field lines and closed white field lines. Helicity is ejected upward along newly open reconnected field lines. At $t = 1500$ the system is slowly relaxing to its final stage.

3.1. Estimators of magnetic helicity conservation

Following Yang et al. (2013), we will compute the volume variation along with the flux of magnetic helicity. From two successive outputs of the studied MHD system, corresponding to two instant τ and τ' , separated by a time interval Δt we will directly compute their respective total helicity $H_V(\tau)$ and $H_V(\tau')$ in the volume \mathcal{V} using the method of Sect. 3.2, and then, the helicity variation rate between these 2 instants: $\Delta H_V/\Delta t = (H_V(\tau') - H_V(\tau))/\Delta t$. We will simultaneously estimate the different sources of fluxes $F_\#$ through the surface of the domain, with $\#$ the different contribution to the total flux in Eq. (30).

We will also time integrate the helicity fluxes at the boundary $H_{\partial V}$:

$$H_{\partial V}(t) = \int_0^t F_{\text{tot}}(\tau) d\tau \quad (39)$$

$$H_{\partial V,\#}(t) = \int_0^t F_\#(\tau) d\tau \quad (40)$$

with $\#$ the different contribution to the total flux of Eq. (30).

One should note that two very different methods are used to compute H_V and $H_{\partial V}$. To derive H_V (and $\Delta H_V/\Delta t$) from Eq. (10), only three components of the magnetic field \mathbf{B} in the whole domain \mathcal{V} are needed: $\mathbf{A}, \mathbf{B}_p, \mathbf{A}_p$ are derived from \mathbf{B} (cf Sect. 3.2). For $H_{\partial V}$ (and F_{tot}), only data along the boundary $\partial \mathcal{V}$ are required. Furthermore, the helicity flux estimation requires the knowledge of the three components of the velocity field on $\partial \mathcal{V}$ (to compute F_{Vn} and F_{Bn}). These quantities are not used to compute H_V . The methods of estimations of H_V and $H_{\partial V}$ are thus fully independent.

A physical quantity is classically said to be conserved when its time variation in a given domain is equal to its flux through the boundary of the domain (cf. Sect. 1). To study the twofold problem of the conservation of relative magnetic helicity and the dissipation of magnetic helicity, we will use two quantities, C_r and C_m , respectively.

For relative magnetic helicity to be perfectly conserved one should have $H_V = H_{\partial V}$, or dH/dt must be equal to F_{tot} , i.e. the volume terms of the relative helicity variation should be null. By estimating, C_r ,

$$C_r = \frac{\Delta H_V}{\Delta t} - F_{\text{tot}} \simeq \left. \frac{dH}{dt} \right|_{\text{diss}} + \left. \frac{dH}{dt} \right|_{\text{Bp,var}}, \quad (41)$$

we can determine the level of conservation of relative magnetic helicity. As already noted in Sect. 2.3 even if \mathbf{B} evolves within

the ideal MHD paradigm, the relative magnetic helicity evolution cannot be written with a classical equation of conservation as the term in $dH/dt|_{\text{Bp,var}}$ is not a flux integral through a boundary and generally does not vanish.

Secondly, we aim to determine the dissipation of the magnetic helicity of the studied field (Eq. 24), i.e. by estimating C_m equal to:

$$C_m = \frac{\Delta H_V}{\Delta t} - F_{\text{tot}} - \left. \frac{\Delta H_V}{\Delta t} \right|_{\text{Bp,var}} \simeq \left. \frac{dH}{dt} \right|_{\text{diss}} \quad (42)$$

Our estimation of C_m is done independently of the estimation of $dH/dt|_{\text{diss}}$. Our method thus does not require the knowledge of the electric field \mathbf{E} , which is a secondary quantity in most MHD problems. The dissipation is thus estimated in a way which is completely independent from the non-ideal process, e.g. the precise way magnetic reconnection is developing.

As the estimators measure both physical helicity variations and numerical errors, we will use different non-dimensional criteria to quantify the level of helicity conservation and the precision of our measurements. We will compute the relative accumulated helicity difference ϵ_H :

$$\epsilon_H(t) = \frac{H_V(t) - H_{\partial V}(t)}{H_{\text{ref}}} \simeq \frac{\int_{\tau=0}^t C_m d\tau}{H_{\text{ref}}} \quad (43)$$

with H_{ref} a normalizing reference helicity, physically significant for the studied system (e.g. the maximum H_V value in the studied interval).

At each instant, C_m expresses the rate of dissipation of helicity, i.e., our numerical estimation of Eq. (24). It can include artefact fluctuations due to the numerical precision of the flux estimation and the time derivation of H_V . In addition, as helicity is a signed quantity both positive and negative helicity can be generated by non-ideal effects. It may be relevant to know the time integrated absolute variation of helicity, as a function of time. Hence, we define another metric, ϵ_{C_m} ,

$$\epsilon_{C_m}(t) = \frac{\int_{\tau=0}^t |C_m| d\tau}{\int_{\tau=0}^t |\Delta H_V/\Delta t| d\tau}, \quad (44)$$

where the absolute values guarantee that we take an upper limit of the dissipation. Along with ϵ_H , ϵ_{C_m} allows us to evaluate the level of dissipation and the numerical precision of our measurements. The measure of low value of ϵ_{C_m} and ϵ_H can thus provide a clear demonstration of the level of conservation of relative

magnetic helicity. For two methods having similar ϵ_H , a higher value of ϵ_{Cm} would indicate measurements presenting larger numerical errors.

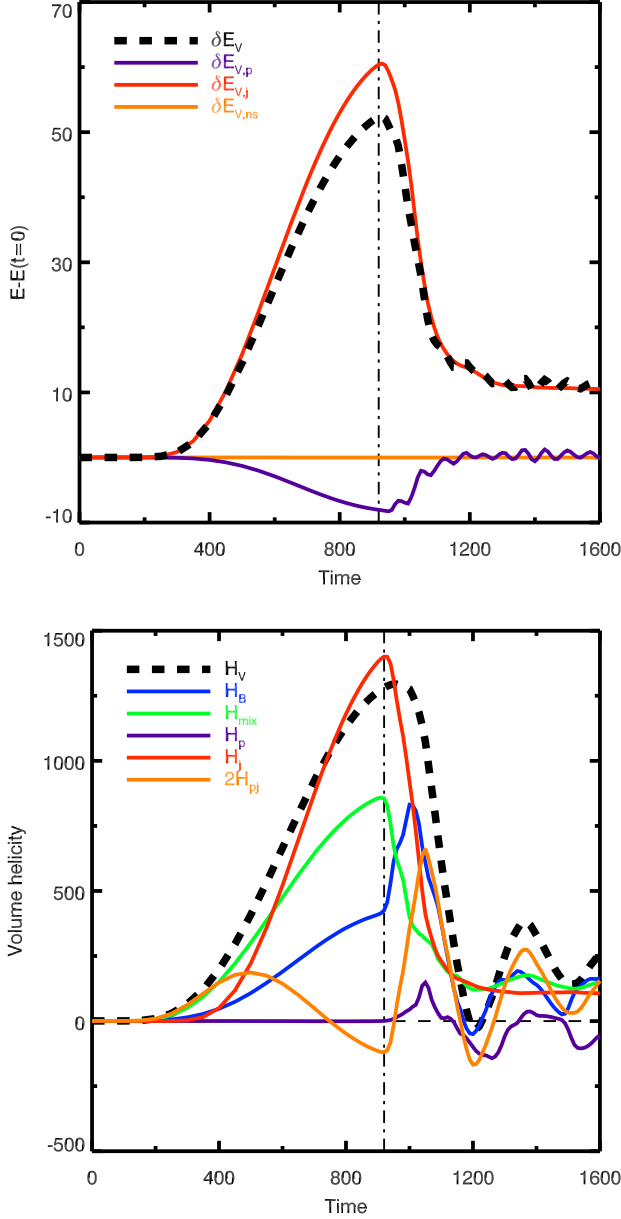


Fig. 2. **Top panel:** Time evolution of the different magnetic energies relative to their respective initial values: total (δE_V , black dashed line), of the potential magnetic fields ($\delta E_{V,p}$, purple line), of the non solenoidal component ($\delta E_{V,ns}$, orange line), and free magnetic energy ($\delta E_{V,j}$, red line). **Bottom panel:** Relative magnetic helicity (H_V , black dashed line) and its decompositions, Eqs. (11,15), computed with the practical DeVore method: proper helicity (H_m , blue line), potential field helicity (H_p , purple line), and the mixed helicity (H_{mix} , green line); current-carrying helicity (H_j , red line), mutual helicity between potential and current-carrying fields (H_{pj} , orange line). In both panels, as in all figures hereafter, the dot-dashed vertical line at $t = 920$ indicates the transition between the quasi-ideal-MHD/energy accumulation and the non-ideal/jet generation phases.

3.2. Volume helicity and flux computations in the DeVore gauge

In order to compute magnetic helicity in the volume, H_V , we will use the method presented in Valori et al. (2012). All the vector potentials will be computed using the gauge presented in DeVore (2000), with no vertical component:

$$A_z = A_{p,z} = 0 \quad (45)$$

Under this assumption, the vector potentials can be computed in the volume using a 1D integral (c.f. Equations (10) or (11) of Valori et al. 2012) with a 2D partial differential equation to be solve at the bottom or top boundary (c.f. Equations (9) or (12) of Valori et al. 2012). We tested that we did not obtained significant differences on the helicity conservation and dissipation properties whether the integration was done from the top or the bottom boundary of the system. In the present manuscript the results were obtained with the integration performed from the top boundary.

While other gauge choices could be explored, choosing the DeVore gauge, Eq. (45), is here motivated by the fact that it is numerically efficient and convenient. This selection leaves a freedom for the gauge of \mathbf{A} and \mathbf{A}_p which can be independent, i.e. not linked as in Eq. (31). In our computation of \mathbf{A}_p and \mathbf{A} , we use gauge freedom to additionally enforce that the two vectors are equal at the top (see Eq. 29 of Valori et al. 2012):

$$\mathbf{A}(z_{top}) = \mathbf{A}_p(z_{top}) \quad (46)$$

We still have a freedom on $\mathbf{A}_p(z_{top})$ as expressed by the 2D partial differential equation (20) of Valori et al. (2012). Here, we select their particular solution expressed by their equations (24,25). With this additional choice, both \mathbf{A}_p and \mathbf{A} are uniquely defined (modulo a constant) by vertical integration starting from the top boundary.

Practically, we first determine the potential field by solving the solution of the Laplace equation (17) for ϕ of Valori et al. (2012). Then we can compute the potential vectors from a direct 1D integration of B_z starting from the top boundary (c.f. Section 3.3 of Valori et al. 2012). We refer to this method below as the "practical DeVore method" since it is efficient and easy to implement.

While the condition (46) holds at the top boundary it does not hold at the other boundaries. Indeed, Eq. 13 in Valori et al. (2012) can be used to show that the difference between the value of \mathbf{A}_p and \mathbf{A} at the bottom boundary is equal to $\hat{\mathbf{z}} \times \int_{z_1}^{z_2} (\mathbf{B} - \mathbf{B}_p) dV$. It is important to note that the conditions of $H_{mix} = 0$ and $\mathbf{A}_p \cdot d\mathbf{S}|_{\partial V} = 0$ are never enforced in this gauge. The relative helicity terms H_{mix} and H_p and the helicity flux term F_{AAp} and F_ϕ can thus never be considered null (see also Section 3.4 in Valori et al. 2012, for more details). Therefore, for the practical DeVore method, the helicity variation of the system is given by the general formula of Eq. (23).

Alternatively, we will determine the helicity assuming a Coulomb gauge for \mathbf{A}_p only. The DeVore and Coulomb gauge are indeed compatible for a potential field. We can solve the 2D partial differential equation (20) of Valori et al. (2012) as a Poisson problem to obtain \mathbf{A}_p (see their Eq. (41), but translated to the top boundary). This implies by construction that \mathbf{A}_p is simultaneously respecting the DeVore, Eq. (45), and the Coulomb gauge, Eq. (33). In the following, we will refer to this method as the "DeVore-Coulomb method".

While still following the DeVore condition (45), \mathbf{A} computed in the DeVore-Coulomb method will be different from \mathbf{A} computed in the practical DeVore method because the boundary condition, Eq. (46) is different. In particular, the distributions of \mathbf{A}

at the bottom boundary will be significantly different with each methods, hence leading to very different values of F_{Bn} and F_{Vn} . In Appendix A, we present an additional test, with \mathbf{A}_p computed in the Coulomb gauge but where \mathbf{A} is not satisfying the condition (46).

3.3. Test data set

In order to test the helicity conservation, we employ a test-case 3D MHD numerical simulation of the generation of a solar coronal jet Pariat et al. (2009). Figure 1 presents snapshots of the evolution of the magnetic field. The simulation assumes an initial uniform coronal plasma with an axisymmetric null point magnetic configuration (left panel). The magnetic null point is created by embedding a vertical dipole below the simulation domain, and adding an uniform volume vertical magnetic field of opposite direction in the domain. The null point present a fan/spine topology, dividing the volume in two domains of connectivity, one closed surrounding the central magnetic polarity, and one open.

The computation are performed with non-dimensional units. We analyze the time evolution of the magnetic field from $t = 0$ to $t = 1600$. The time steps between two outputs is $\Delta t = 50$ for $t \in [0; 700]$, during the accumulation phase, and, $\Delta t = 10$ for $t \in [700; 1600]$ during the dynamic phase of the jet. The analysis volume \mathcal{V} is a subdomain of the larger discretized volume employed in the original MHD simulation: $[-6, 6]$ in x -, $[-6, 6]$ in y -, and $[0, 12]$ in the z -directions, thus only taking into account the region of higher resolution (c.f. Figure 1 of Pariat et al. 2009).

The ideal MHD equations are solved with the ARMS code based on Flux Corrected Transport algorithms (FCT, DeVore 1991). The parallelisation of the code is ensured with the PARAMESH toolkit (MacNeice et al. 2000). In the original simulation of Pariat et al. (2009), reconnection is strongly localized thanks to the use of adaptive mesh refinement methods at the location of the formation of the thin current sheets involved in reconnection (c.f. Appendix of Karpen et al. 2012). In the present paper in order to keep the resolution of the domain constant we have however switched off adaptivity, so the resolution is equal to the initial one (as in Figure 1 of Pariat et al. 2009), throughout the whole simulation.

The variation of energies relatively to their initial value, $\delta E_{\mathcal{V}}(t) = E_{\mathcal{V}}(t) - E_{\mathcal{V}}(t = 0)$ is displayed in Fig. 2, top panel. The index \mathcal{V} indicates that the energy is computed by a volume integration. The total energy, $E_{\mathcal{V}}$, in the domain can be decomposed as:

$$E = E_{\mathcal{V},p} + E_{\mathcal{V},j} + E_{\mathcal{V},ns}, \quad (47)$$

where $E_{\mathcal{V},p}$ and $E_{\mathcal{V},j}$ are the energies associated to the potential and current-carrying solenoidal contributions and $E_{\mathcal{V},ns}$ is the sum of the nonsolenoidal contributions (see Eqs. (7,8) in Valori et al. 2013, for the corresponding expressions). Initially the system is fully potential and $E_{\mathcal{V}} = E_{\mathcal{V},p}$. Energy is injected in the system by line-tied twisting motions of the central polarity. The axisymmetric boundary motions are preserving the distribution of B_z at the bottom boundary. Magnetic free energy and helicity accumulates monotonically increasing the twist in the closed domain (Fig. 1, central left panel). The potential field energy $E_{\mathcal{V},p}$ decreases slightly during the accumulation phase because of the bulge of the central domain which is changing the distribution of the field on the side and top boundaries. The $E_{\mathcal{V},ns}$ term is almost constantly null, an indication of the excellent solenoidality of the system

Eventually, around $t \simeq 920$, the system becomes violently unstable: magnetic reconnection sets in and the closed twisted field lines reconnect with the outer open field lines. A steep decrease of the free magnetic energy is observed in Fig. 2: 83% of the maximum free magnetic energy is dissipated/ejected/transformed. Through reconnection, twist and helicity are expelled from the central domain, inducing a large scale kink wave which exits through the top boundary (Fig. 1, central right panel). This large scale non-linear magnetic wave, simultaneously compresses and heats the plasma, inducing the generation of an untwisting jet that can observationally be interpreted as a blowout jet (Patsourakos et al. 2008; Pariat et al. 2015). The driving motions had been slowly ramped down at the time of the trigger of the jet so that few energy and helicity are injected from the lower boundary after the jet onset (cf. Figure 6 of Pariat et al. 2009). In the final stage the system slowly relaxes to a configuration similar to its initial state, with the potential field energy being similar to its initial value, and only few field lines remaining twisted, next to the inversion line (Fig. 1, right panel).

This simulation thus presents two distinct phases typical of active events: before $t \simeq 920$, a phase with a slow ideal accumulation of magnetic helicity and energy, and after $t \simeq 920$, an eruptive phase of fast energy dissipation and helicity transfer involving non-ideal effects. In the first phase, the system behaves very close to ideality as demonstrated in a benchmark with a strictly ideal simulation (Rachmeler et al. 2010). In the non-ideal phase, Pariat et al. (2009) showed that 90% of the helicity was eventually ejected through the top boundary by the jet, and Pariat et al. (2010) showed the high reconnection rate processing the magnetic flux during the jet. These two phases allow us to test the conservation of helicity in two very distinct paradigms of MHD.

In the following we test the helicity conservation with the above MHD simulation. We first use the practical DeVore method (see Sect. 4). We next test the effect of the gauge choice on our results, by using the DeVore-Coulomb method (see Sect. 5). We will show how the gauge choice possibly affects the evolution of each terms.

4. Magnetic helicity conservation in the practical DeVore gauge

Our first study of the helicity variations is performed in the practical DeVore gauge. The only assumptions on the potential field and the vector potentials are given by Equations (9), (45) & (46). This corresponds to a general case where the helicity variation of the system is provided by Eq. (23).

4.1. Helicity evolution

Figure 2, bottom panel, presents the evolution of the relative magnetic helicity $H_{\mathcal{V}}$ in the system. Similarly to magnetic energy the two phases of the evolution are clearly marked. The first phase corresponds to a steady accumulation of magnetic helicity, while the second corresponds to the blowout jet. The latter is associated with a steep decrease of magnetic helicity. As the system relaxes the magnetic helicity value oscillates. These oscillations are related to the presence of a large scale Alfvénic wave which is slowly damped after the jet. This oscillations can also be seen in the total energy but with a much smaller relative amplitude (so small that $E_{\mathcal{V}}$ still decreases monotonically).

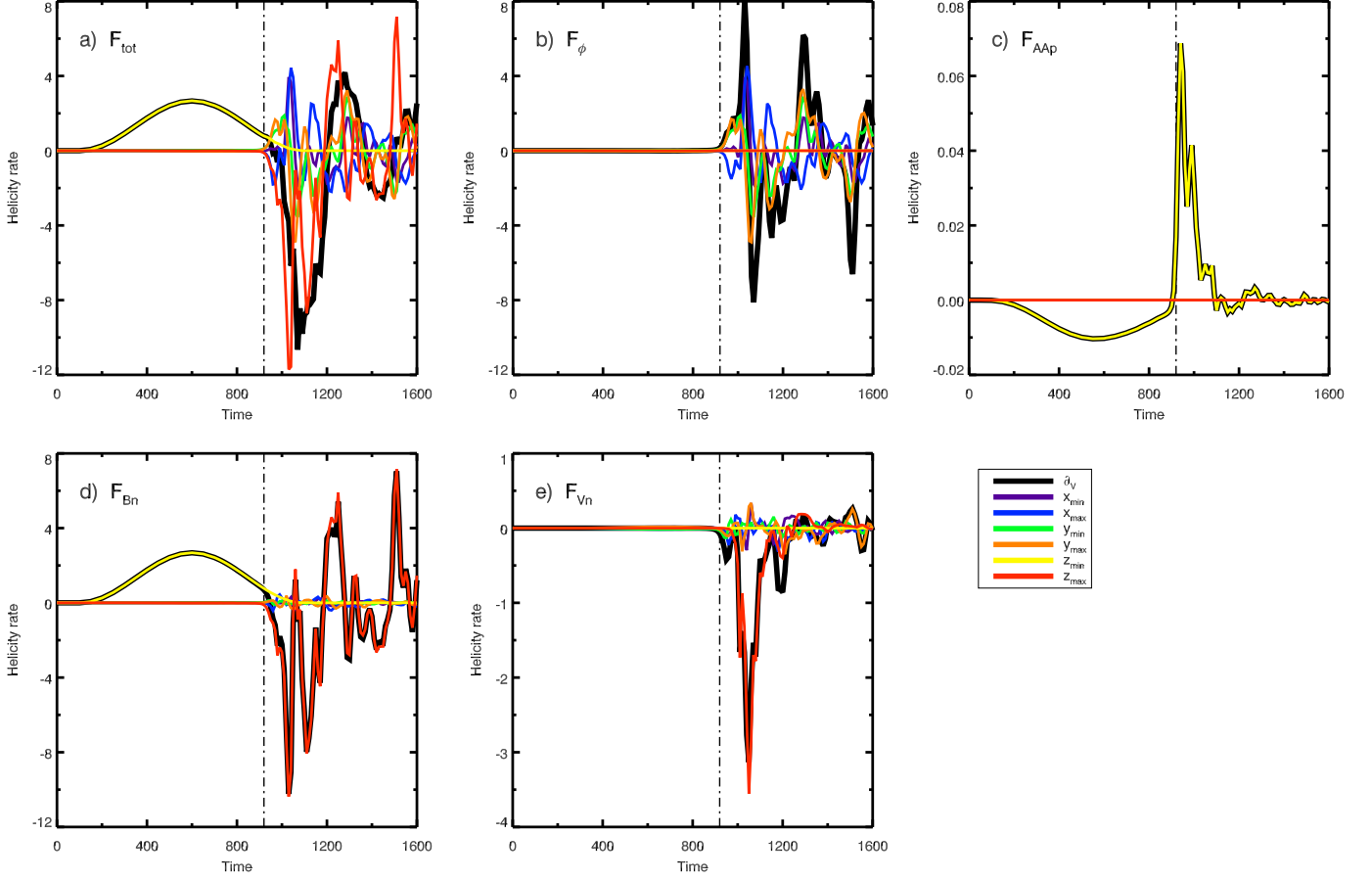


Fig. 4. Total helicity flux, F_{tot} (Eq. 30), and the the terms composing it, Eqs. (26 - 29), through the different boundaries computed in the practical DeVore Gauge. (a): F_{tot} ; (b): F_{ϕ} ; (c): F_{AAp} ; (d): F_{Bn} ; (e): F_{Vn} . In each plot the dark line correspond to the sum of the flux through all the boundaries while the color lines correspond to a flux through a particular boundary (purple and blue: left and right x sides, green and orange: front and back y sides; yellow and red: bottom and top z sides).

The right panel of Fig. 2 also presents the decomposition of the relative magnetic helicity, H_V , in H_m , H_{mix} and H_p of Eqs. (11 - 13). During the accumulation phase, H_V is dominated by H_{mix} while H_p is almost constantly null. During the jet, one can see strong fluctuations of the relative importance of these terms. H_m , H_{mix} and H_p eventually present contributions of similar amplitude. However, because these terms are not gauge invariant, their value in a different gauge might be quite different (cf. Sect. 5).

The decomposition of H_V with H_j and H_{pj} of Eqs. (15 - 17) is also plotted in Fig. 2, bottom panel. H_j has an evolution comparable to $\delta E_{V,j}$. It captures most of the helicity evolution both during the accumulation and jet phases. In contrast, the mutual helicity between the potential and the current carrying fields contains mostly oscillations. Therefore, H_j , which is a gauge invariant quantity, is a promising quantity to analyze during a jet/eruption.

4.2. Helicity fluxes

We compute the time variations $\Delta H_V / \Delta t$ of the helicity determined with the volume integration method, and compare it with the different terms of the relative magnetic helicity flux through the whole system boundary (Fig. 3). This shows that the helicity variation $\Delta H_V / \Delta t$ is very closely matching the curve of F_{tot} , indicating that the variation of helicity in the domain are tightly

related to the flux of relative helicity through the boundary. The core results of this study is that indeed magnetic helicity is very well conserved in the studied simulation, both during the quasi-ideal-MHD and non-ideal phases.

During the non-ideal/jet phase, strong fluctuations are observed for all terms but the F_{AAp} term. The later is constantly negligible relatively to the others. On the other hand we see that the F_{ϕ} term displays important fluctuations, frequently of similar amplitude and opposite sign to F_{Bn} . The F_{ϕ} term clearly cannot be neglected in this gauge.

The analysis of the flux of each terms through each individual boundary is insightful (Fig. 4), though one must bear in mind that the plotted terms are not gauge invariants. Although its amplitude is extremely small, a finite flux of F_{AAp} is present only at the bottom boundary during the whole evolution of the system: there is no flux on the sides because of the DeVore gauge (Eq. 45), and no flux on the top because of the imposed condition of Eq. (46).

During the ideal phase, the flux of helicity is completely dominated by the F_{Bn} term originating from the bottom photospheric boundary. This is consistent with the fact that the system is indeed driven by horizontal shearing motions at the bottom boundary. No remarkable helicity flux is observed at the other boundaries during this period. Helicity is thus accumulating in the volume V .

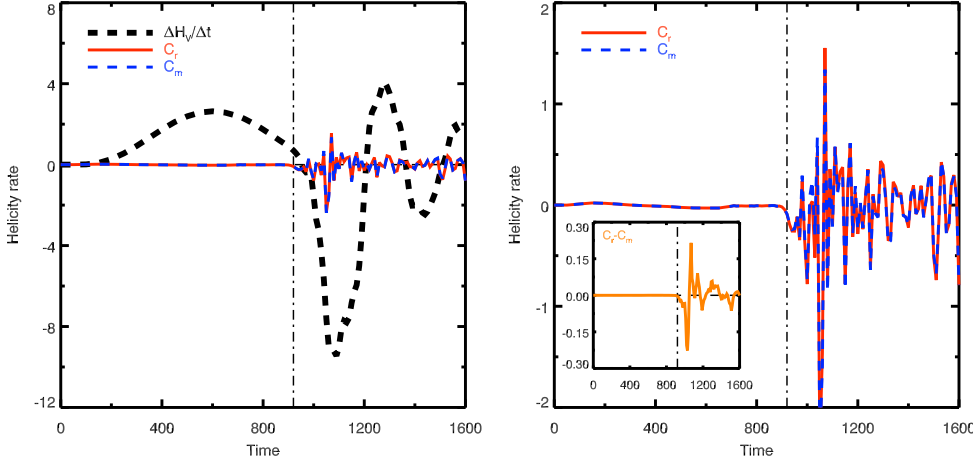


Fig. 5. Difference between the helicity variation rate and the boundary helicity flux, computed in the practical DeVore gauge. The relative helicity conservation criterion C_r (red line, Eq. 41) and the magnetic helicity dissipation criterion C_m (blue line, Eq. 42) are plotted relatively to the helicity variation rate ($\Delta H_V/\Delta t$, black dashed line, left panel) and in their own amplitude range (right panel). The insert in the right panel present the difference between C_r and C_m , i.e. the potential helicity volume variations $dH/dt|_{p,var}$.

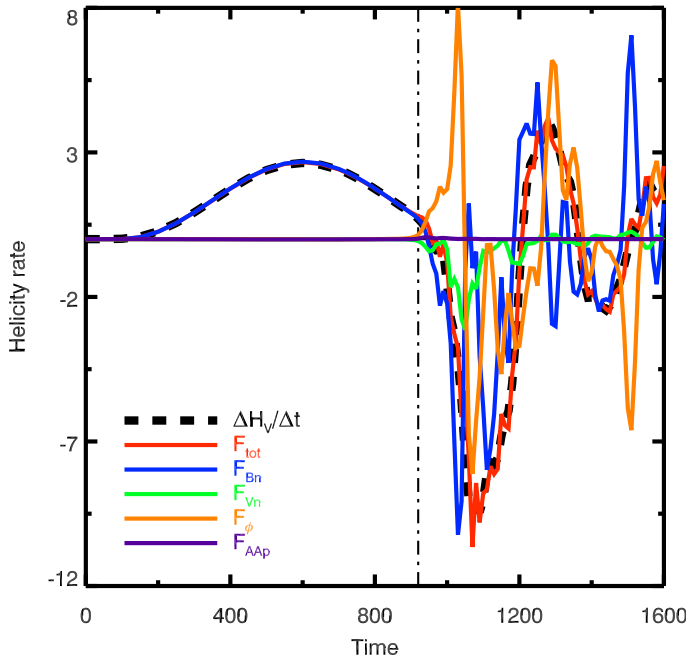


Fig. 3. Comparison of the helicity variation rate and the helicity flux computed in the practical DeVore gauge. The helicity variation rate ($\Delta H_V/\Delta t$, dashed line) is derived from the volume integration method. The total helicity fluxes through the whole surface, ∂V , of the domain is F_{tot} (red line). This flux is composed of, Eqs. (26 - 29): F_{Vn} (green line), F_{Bn} (blue line), F_{ϕ} (orange line), F_{AAp} (purple line).

During and after the jet ($t > 920$) important helicity fluxes are noted in the side and top boundaries, while the bottom flux is now negligible as the boundary flows have been ramped down in amplitude. A large flux of helicity occurs at the top boundary (red curves), dominated by the F_{Bn} term and to a lower extent by the F_{Vn} term. This corresponds to the ejection of helicity by the jet, driven by a large-scale non-linear torsional wave (Pariat et al. 2009, 2015). F_{Vn} is peaked at the time of the passage of the bulk of the jet through the top boundary. The flux of F_{Bn} and F_{Vn} through the side boundaries, while present, is comparatively small. However, the side boundaries see the transit of important flux F_{ϕ} . No specific side boundary is dominating the total value of F_{ϕ} . Due to the DeVore gauge (Eq. 45), F_{ϕ} is null at the bottom and top boundaries.

We conclude that, computed with the particular DeVore method, the total flux of helicity F_{tot} during the jet consists of complex transfer of helicity through all the boundaries of the system.

4.3. Helicity conservation

In order to better estimate the helicity conservation and dissipation, we have plotted in Fig. 5 the criterions C_r and C_m (see their definition in Sect. 3.1). We observe that these two criterions are almost equal. They differ only by the term describing the volume variation of the potential helicity $dH/dt|_{BP,var}$ of Eq. (26). Our calculation finds that $dH/dt|_{BP,var}$ is extremely small compared to the variation of C_r and C_m (see inset in the right panel of Fig. 5), even though we are not explicitly enforcing $\nabla \cdot \mathbf{A}_p = 0$. This term is smaller than the combined effect of the real helicity dissipation $dH/dt|_{diss}$ and the numerical errors on the volume and flux helicity measurements.

The curve of C_r demonstrates that magnetic helicity is extremely well conserved during the ideal-MHD phase of the simulation and that it is also very well conserved during the non-ideal phase. Moreover, during the ideal phase the amplitude of $|C_m|$ does not exceed 0.029 which is 1% of the maximum amplitude of helicity variation during the period. At the end of this period, one also have $\epsilon_{Cm}(t = 920) < 1\%$, thus helicity is very weakly dissipated, as theoretically expected (Woltjer 1958).

During the non-ideal phase, C_r and C_m show high frequency fluctuations around the null value, decorrelated from the fluctuation of helicity in the system. Our analysis indicates that while these fluctuations could originate from the real physical term $dH/dt|_{diss}$, they are in fact dominated by the numerical precision on the estimation of F_{tot} . From the same simulations data, we will see in Sect. 5 that the computation with the DeVore-Coulomb method reduces these fluctuations.

In Fig. 6 we have plotted the variation of magnetic energy and helicity in the system computed with a volume integration and from the integration of the Poynting and helicity fluxes through the boundaries of the system. During the ideal-MHD phase both magnetic energy and helicity are well conserved, their volume variations being equal to their boundary fluxes. During the non-ideal phase, while magnetic helicity is still very well conserved magnetic energy is clearly not. When the jet is generated, the magnetic energy quickly decreases: part of it is ejected through the top boundary by the jet, but for most part it is dissipated in the reconnection current sheet and transformed

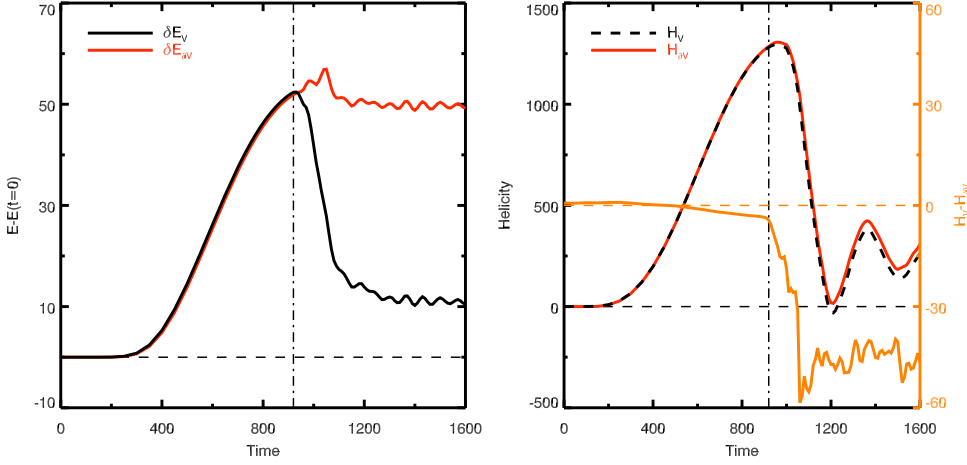


Fig. 6. **Left panel:** Comparison of the evolution of magnetic energy (δE_V , black line) in the volume with the time integration of the Poynting flux through the whole surface of the domain ($\delta E_{\partial V}$, red line). **Right panel:** Comparison of the evolution of helicity obtained by volume integration (H_V , black dashed line) with the time integration of the helicity flux through the whole surface of the domain ($H_{\partial V}$, red line), computed with the practical DeVore gauge. Their difference is plotted in orange on a different range of amplitude (c.f. right axis).

in other forms of energy. When the simulation is stopped, we determined that about 17% of the magnetic energy injected in the system by the bottom boundary motions remains in the system, 21% is directly ejected with the jet and the 62% remaining are dissipated/transformed in other form of energy (see also Pariat et al. 2009). As expected during reconnection, magnetic energy is strongly non-conserved.

The situation for magnetic helicity is very different. At the end of the ideal phase the maximum difference between H_V and $H_{\partial V}$ is of 3.5 units. During this period a total of $H_{ref} = \max(H) = 1265$ units have been injected and only a fraction $\epsilon_H(t = 920) = 0.3\%$ is lost. The ideality of the system is thus very well maintained. During the non-ideal phase, the maximum difference between H_V and $H_{\partial V}$ is equal to 58 units (orange line, Fig. 6, right panel) and $\epsilon_H(t = 1600) = 4.5\%$. The jet is able to transport away a huge fraction, 77% of the helicity of the system. Finally, about 19% of the helicity remains in the system at $t = 1600$ while H_V is still oscillating slightly (Fig. 6, right panel).

Summarizing, with a generic gauge, for a solar-like active events, we have thus confirmed Taylor (1974) hypothesis that magnetic helicity is very well conserved even when non-ideal processes are acting. The relative helicity dissipation is 15 times smaller than the relative magnetic energy dissipation. However, with the practical DeVore Gauge, we observed that C_m is strongly fluctuating, possibly limited by numerical precision. We now want to test whether we can obtain better results with a different gauge, simultaneously allowing us to test the gauge-invariance of the different helicity variation terms.

5. Magnetic helicity conservation in the DeVore-Coulomb case

Relative magnetic helicity has been defined as a gauge-invariant quantity (see Sect. 2.3). However, the surface flux of relative helicity is not, neither are the individual terms that defines the flux. In the following we will study the influence of computing magnetic helicity using a different gauge. The vector potential of the potential field, \mathbf{A}_p , is now computed in the DeVore-Coulomb gauge, but not \mathbf{A} (since the Coulomb gauge is only compatible with DeVore gauge for a potential field). Because we also impose the condition of Eq. (46), \mathbf{A} is however also recomputed during our DeVore-Coulomb method (see Sect. 3.2). The case where Eq. (46) is not enforced is briefly discussed in Sect. A

5.1. Gauge dependance

In Fig. 7, we have plotted the evolution of the relative magnetic helicity, H_V in the DeVore-Coulomb gauge for \mathbf{A}_p . The left panel shows the comparison with the same quantity computed using the practical DeVore gauge, used in the previous section. We observe that the two curves matches almost perfectly. Their maximum differences is at most 3.5 units, and their maximum relative helicity difference is of 0.3%. The gauge invariance is thus very well respected for our estimation of the relative magnetic helicity.

In Fig. 7, right panel, the decomposition of the relative magnetic helicity is plotted. Unlike with the gauges of the practical DeVore method (cf. Fig. 2), H_p is now almost constantly null. It is not strictly null since we are not imposing $\mathbf{A}_p \cdot d\mathbf{S}|_{\partial V} = 0$ on the side boundaries. We see that when the jet is developing H_p fluctuates slightly (barely visible on Fig. 7). The helicity is thus distributed differently between H_m , H_{mix} and H_p in the DeVore-Coulomb case compared to the practical DeVore gauge. As expected these helicity terms are indeed not gauge invariant. Most of the helicity that was carried by H_p for the practical DeVore gauge is now carried by H_m while H_{mix} remains very similar in both gauges (compare the green curves in Figs. 2 and 7). In general, we have to expect a completely different distribution of H_m , H_{mix} and H_p in other gauges. On the other hand, the quantity H_j and H_{pj} remains equal in both gauge computation (with the same precision as H_V). This is expected since H_j and H_{pj} are gauge-invariant quantities.

For the helicity fluxes, in the DeVore-Coulomb case the time variations $\Delta H_V / \Delta t$ of the helicity follows tightly the helicity flux through the boundary (Fig. 8, left panel). While $\Delta H_V / \Delta t$ is gauge invariant, the different contribution of F_{tot} are not. By comparison with Fig. 3 one observes that F_{Bn} , F_{Vn} , F_ϕ have a very different evolution. While F_ϕ was presenting fluctuations of the same amplitude as F_{tot} in the practical DeVore gauge, this term is now very weak throughout the evolution of the system with the DeVore-Coulomb method. The F_{Bn} term is dominating F_{tot} in the DeVore-Coulomb case both during the ideal MHD phase (at the bottom boundary) and during the non-ideal period (at the top boundary). F_{Vn} is weaker than with the practical DeVore gauge. The F_{AAp} term remains negligible in both cases, although with a different choice of gauge, F_{AAp} can contribute significantly to F_{tot} (see Appendix A).

The impact of the gauge dependance is more strikingly illustrated when looking at the time integrated helicity fluxes $H_{\partial V, \#}$ (Eq. 40) through the boundaries. Their evolution is presented

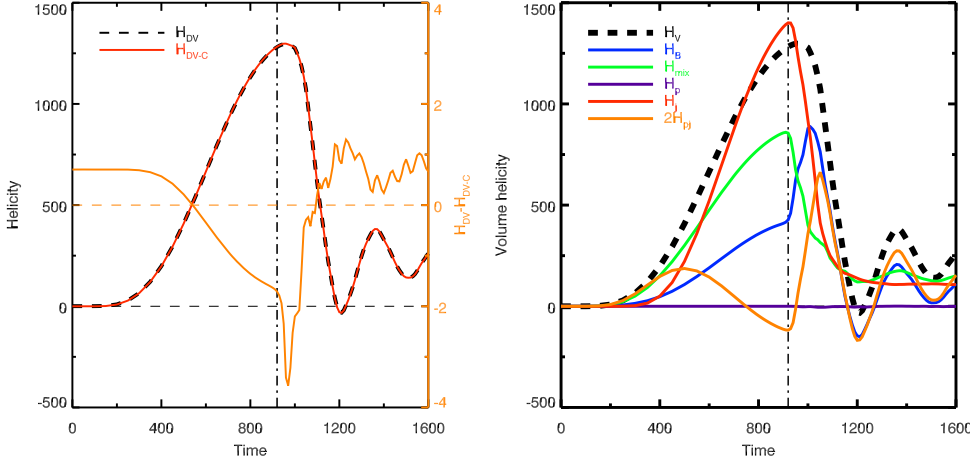


Fig. 7. **Left panel:** Evolution of the magnetic helicity computed with the practical DeVore method (dashed line) and with the DeVore-Coulomb method (red line). Their difference is plotted with an orange line (c.f. right axis). **Right panel:** Volume magnetic helicity (H_V , black dashed line) and its decomposition in the DeVore-Coulomb case. The plotted curves are similar to Figure 2, bottom panel.

in Fig. 9 for the two gauges computations. While H_V remains gauge invariant (Fig. 7, left panel), as expected the $H_{\partial V, \#}$ present different profiles in the different gauges. $H_{\partial V, \phi}$ is small when computed with the DeVore-Coulomb gauge while it presents large amplitudes in the practical DeVore gauge. $H_{\partial V, V_n}$ and $H_{\partial V, B_n}$ present smaller mean absolute values when \mathbf{A}_p follows the DeVore-Coulomb conditions. Their ratio is also strongly dependent on the set of gauges employed.

Numerous studies have computed $H_{\partial V, V_n}$ & $H_{\partial V, B_n}$ and followed their time evolution in observed active regions (e.g. Zhang et al. 2012; Liu et al. 2013, 2014b,a). These terms have been incorrectly called "emergence" and "shear" terms and physical insight have been tried to be extracted from their values and respective ratio. However, as these terms are not gauge invariant, one must question the pertinence of such results. $H_{\partial V, V_n}$ and $H_{\partial V, B_n}$ cannot be studied independently as, for a given \mathbf{v} , their intensities and respective values can be simply modified by a change of gauge. This extends the conclusion of Démoulin & Berger (2003) which showed that only the sum of $H_{\partial V, V_n}$ and $H_{\partial V, B_n}$ can be derived when only the tangential velocity components are known. More generally, the study presented here shows that, even when the full velocity field on the boundary is known, $H_{\partial V, V_n}$ and $H_{\partial V, B_n}$ can present different values depending on the gauge used for the computation. In summary, only the sum $H_{\partial V}$ of all the flux terms, $H_{\partial V, \#}$, is carrying a meaningful information.

5.2. Helicity dissipation

The estimators C_r and C_m (cf. Sect. 3.1) enable us to study the helicity conservation and dissipation. Since $dH/dt|_{B_p, var}$ is null because of the Coulomb condition on the potential field, C_m and C_r are equal to our numerical precision. There is no volume helicity variation due to the potential field. The conservation of the relative helicity is only limited by the actual dissipation helicity $dH/dt|_{diss}$. The curve of C_r for the DeVore-Coulomb computation confirms that magnetic helicity is gauge-invariantly well conserved during the simulation (Fig. 8, right panel).

In fact the computation in the DeVore-Coulomb case even improves our estimation of the helicity dissipation. During the quasi-ideal phase, the computation of C_m in the DeVore-Coulomb case is equal to its value with the practical DeVore case. However we observe that in the non-ideal phase C_m presents smaller oscillations when computed in the DeVore-Coulomb case. Theoretically C_m , being equal to the helicity dis-

sipation, $dH/dt|_{diss}$, should be gauge invariant. However we observe that C_m can change by a factor 2 when computed with the different methods. With the practical DeVore method, the fluctuations of C_m peak at 25% of the maximum amplitude of helicity variation during the non-ideal phase, while the peak is only equal to 4.5% of the amplitude in the DeVore-Coulomb case. The non-dimensional criterion, ϵ_{C_m} at the end of the simulation is equal to 7% in the practical DeVore case while it is limited at 2.7% with the DeVore-Coulomb method (see the inset in Fig. 8, right panel).

While the measure of H_V is done with a high precision (see above), the estimation of the helicity flux, F_{tot} induces more numerical errors. The high frequency oscillations observed in the different terms contributing to F_{tot} are the symptom of the lower level of precision. The criterion C_m is in fact equal to the helicity dissipation plus the numerical errors in the volume helicity variations and in the helicity flux. What C_m is in fact providing is an upper value for the helicity dissipation. In the DeVore-Coulomb case, C_m is likely providing a smaller value, hence a better constraint on $dH/dt|_{diss}$, because the F_{tot} is dominated by the errors on F_{Bn} , while in the practical DeVore case, the errors of the fluctuating F_{Bn} , F_{Vn} and F_{ϕ} will sum up with comparable magnitudes.

The measure of the helicity flux through the boundary is the limiting factor that does not permit us to reach numerical precision during the non-ideal phase. It is likely that a higher spatial resolution sampling of the velocity field on the boundary sides will further improve the helicity dissipation estimation. Our present computation of C_m is therefore only an upper bound on the real helicity dissipation $dH/dt|_{diss}$. Because of its lower ϵ_{C_m} , the DeVore-Coulomb method allows us to better bound the helicity dissipation.

In DeVore-Coulomb computation, during the ideal-MHD phase, the maximum difference between H_V and $H_{\partial V}$ is of now of 4.3 units and $\epsilon_H(t = 920) = 0.3\%$. As theoretically expected, the dissipation of magnetic helicity is extremely small when only ideal process are present. It actually demonstrates that the ideality of the system is very well maintained by the numerical scheme during that phase. At the end of the non-ideal phase, the maximum difference between H_V and $H_{\partial V}$ is equal to 28.2 units (orange line, Fig. 9, right panel) and the relative amount of helicity dissipated is smaller than $\epsilon_H(t = 1600) = 2.2\%$.

In absolute value, the dissipation of magnetic helicity is thus very small compare to the magnetic helicity which is ejected away. The dissipated helicity represents less than 3% of the he-

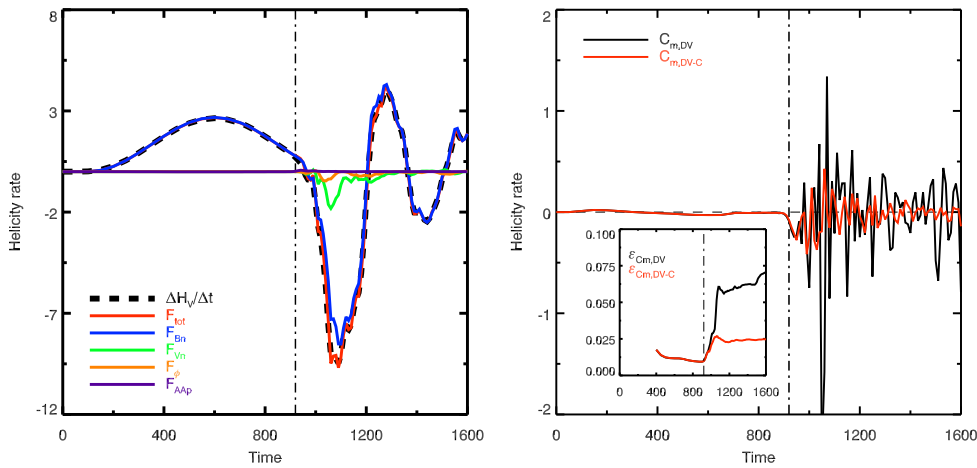


Fig. 8. **Left panel:** Comparison of the helicity variation rate and the helicity flux integrated through the six boundaries of the domain in the DeVore-Coulomb case. The plotted curves are similar than in Figure 3. **Right panel:** magnetic helicity dissipation criterion C_m , Eq. (42), plotted for the practical DeVore (black line) and the DeVore-Coulomb (red line) methods. The insert shows ϵ_{Cm} , Eq. (44), with the same color convention.

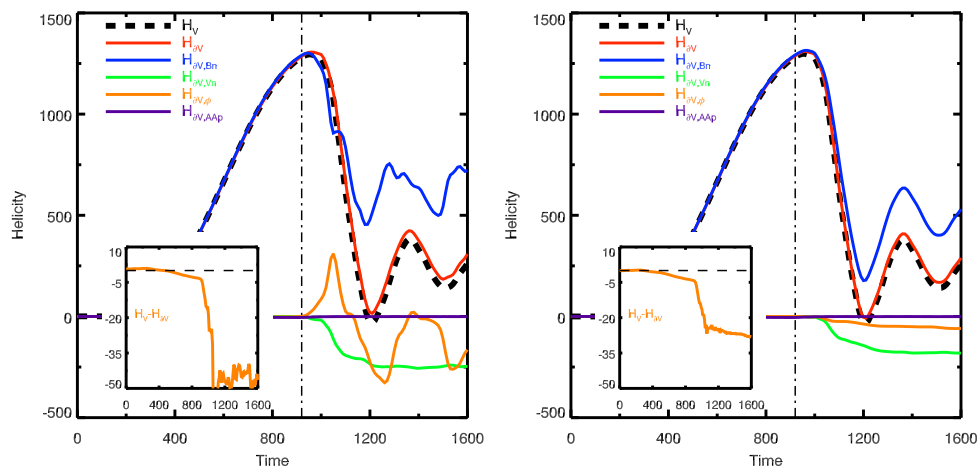


Fig. 9. H_V and $H_{DV,\#}$ evolution in the system computed with different gauge. The left panel has been computed with the practical DeVore method while the right panel has been derived with the DeVore-Coulomb method. The helicity variation (H_V , black dashed line) is derived from the volume integration method. The time integrated helicity is H_{DV} (red line) which can be decomposed in the contribution of its different terms: $H_{DV,Bn}$ (blue line), $H_{DV,Vn}$ (green line), $H_{DV,\phi}$ (orange line) and $H_{DV,AAp}$ (purple line). The orange line in the inserts correspond to the difference between H_V and H_{DV} .

helicity carried away by the jet. If one want to track the helicity evolution of this jet system, the helicity dissipation only represent a very minor contribution, more than one order of magnitude smaller than the helicity staying in the system and the ejected helicity.

This low helicity dissipation is also to be compared to the relative decrease of magnetic energy of 62% which is developing simultaneously. The reconnections generating the jets are thus dissipating/transforming magnetic energy $\gtrsim 30$ times more efficiently than magnetic helicity is dissipated. While vast amounts of energy are lost, magnetic helicity is barely affected by the non-ideal MHD processes at play.

6. Conclusion

Based on the property that magnetic helicity presents an inverse cascade from small to large scale, Taylor (1974) conjectured that magnetic helicity, similarly to pure ideal MHD, is also effectively conserved when non-ideal processes are present. Because of the inherent difficulties to measure magnetic helicity, tests of this conjecture have so far been very limited (cf. Sect. 1). The theoretical development of relative magnetic helicity (Berger & Field 1984), as well as the publication of recent methods to actually measure relative magnetic helicity in general 3D datasets (Thalmann et al. 2011; Valori et al. 2012; Yang et al. 2013) are now opening direct ways to test the conservation of magnetic helicity.

In this manuscript, we have performed the first precise and thorough test of Taylor (1974) hypothesis in a numerical simulation of an active solar-like event: the impulsive generation of a blowout jet (Parat et al. 2009). Following Yang et al. (2013), our methods to test the level of magnetic helicity dissipation relies on the comparison of the variation of relative helicity in the domain with the fluxes of helicity through the boundaries.

This lead us to revise the formulation of the time variation of relative magnetic helicity in a fully bounded volume (cf. Sect. 2.3). As relative magnetic helicity relies on magnetic vector potential, the question of the gauge is a central problematic of any magnetic-helicity related quantity. A general decomposition of the gauge-invariant time variation of magnetic helicity is given in Eq. (23), with no assumption made on the gauges of the magnetic field and of the reference potential field. Furthermore, we discussed how specific gauges and combination of gauges can simplify the formulation of the helicity variation (Sect. 2.4).

Following (Valori et al. 2012), we computed the variation of relative magnetic helicity using different gauges, all of them based on the gauge of Eq. (45) suggested by DeVore (2000). We have been able to test the gauge dependance of several terms entering in the decomposition of relative magnetic helicity and its time variation. We demonstrated that all the quantities that were theoretically gauge-invariant (H , H_j , H_{pj} , dH/dt) were indeed invariant with a very good numerical precision ($< 0.3\%$ of relative error).

Additionally, our analysis showed the effect of using different gauges on gauge-dependant quantities. Of particular interest are the results that we obtain on the F_{Vn} and F_{Bn} terms, Eqs. (26,27), entering in the decomposition of the helicity flux. In some studies (e.g. Liu et al. 2014b,a), these terms have been used to putatively track the helicity contribution of vertical and horizontal plasma flows. However, our computations illustrates that these fluxes (and their ratio) vary with the gauges used hence precluding any meaningful physical insight of their interpretation in term of helicity. Following Démoulin & Berger (2003), we have concluded that only the total helicity flux, F_{tot} , conveys a physical meaning.

Unlike magnetic helicity, even in ideal MHD, we showed that relative magnetic helicity cannot be expressed in general in classical conserved form. Only when computing the reference potential field with the Coulomb gauge, can the variation of the relative magnetic helicity be expressed in ideal MHD as a pure surface flux. Generally, relative magnetic helicity is therefore not a conserved quantity in a classical sense.

In the first phase of our simulation, because of a topological constraint (Pariat et al. 2009), the system is believed to follow tightly an ideal MHD evolution. De facto, in our practical numerical case, we observed during that phase, that relative magnetic helicity is very well conserved, its variation following the time-accumulated flux of helicity with a relative accuracy of 0.3%. The relative helicity dissipation that we obtain is one order of magnitude smaller than the one estimated during the ideal evolution of the simulation tested in Yang et al. (2013). The measure of the helicity dissipation can appear as a practical way to test the level of ideality in a simulation. During the ideal MHD phase, as expected, the Flux Corrected Transport scheme (DeVore 1991) that is used to produce our test numerical simulation, is effectively able to ensure a quasi-ideal evolution with a measurable high degree of precision.

Furthermore, the term-by-term study of the helicity variation enables to determine the real, gauge-invariant, dissipation of the magnetic helicity of the studied magnetic field dH/dt_{diss} , Eq. (24). For a solar-like active event, we have confirmed Taylor (1974) hypothesis that magnetic helicity is very well conserved even when non-ideal processes are acting. For the specific event that we have studied, even when intense magnetic reconnection is present, less than 2.2% of the injected helicity is dissipated. While this is one order of magnitude larger than during the ideal phase, the dissipation of magnetic helicity is more than 30 times smaller than the dissipation of magnetic energy during the same period.

Yang et al. (2013) and this study are paving the way for future, more complete and more extensive tests of the Taylor's conjecture. In parallel to the exploration of the properties of helicity, our study offers also more numerically-oriented applications. In a previous work (Valori et al. 2013) we introduced a diagnostic for numerical discretization of magnetic field dataset that present finite error of non-solenoidality which are impacting the estimation of their magnetic energy. Further studies on the effect of time and spatial resolution, on a wider range of processes and dynamical MHD evolution are now needed. If precise quantitative bounds are placed on the level of helicity dissipation, magnetic helicity will eventually become useful to test numerical MHD codes: the level of helicity dissipation could be used as a quantitative criterion of the quality of numerical experiments over the entire simulated evolution. With respect to the instantaneous divergence metric of the magnetic field, helicity is a complementary, extremely sensitive proxy suitable for testing integral conservation properties. The method of the measure of

the helicity dissipation presented in this manuscript, would open up a new way to benchmark numerical codes.

The strong physical insight that can be gained by studying magnetic helicity is also further raised by our present study. More than forty years after, our numerically precise tests of Taylor (1974) conjecture on a solar like events, confirms that magnetic helicity is a quasi-conserved quantity of MHD. The application of the conservation of magnetic helicity is full of potential for the study of complex natural and experimental magnetised plasma systems. Because of its conservation, magnetic helicity may be the *raison d'être* of the existence of coronal mass ejection (Rust 1994; Low 1996). Magnetic helicity can be tracked to characterise and relate the evolution of coronal active region with interplanetary magnetic clouds (e.g. Mandrini et al. 2005; Nakwacki et al. 2011).

Finally, the impact of magnetic helicity conservation on the magnetic reconnection mechanism remains to be better understood (Russell et al. 2015). While it has been observed that magnetic helicity can significantly modify the reconnection dynamics (Linton et al. 2001; Del Sordo et al. 2010), how magnetic helicity is actually redistributed in the system, at quasi-constant total value by magnetic reconnection, still needs to be determined.

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References

- Alexakis, A., Mininni, P. D., & Pouquet, A. 2006, *The Astrophysical Journal*, 640, 335
- Antiochos, S. K. & DeVore, C. R. 1999, *Magnetic Helicity in Space and Laboratory Plasmas: Geophysical Monograph* 111. Edited by Michael R. Brown, 187
- Barnes, C. W., Fernández, J. C., Henins, I., et al. 1986, *Physics of Fluids*, 29, 3415
- Berger, M. A. 1984, *Geophysical and Astrophysical Fluid Dynamics* (ISSN 0309-1929), 30, 79
- Berger, M. A. 2003, *Advances in Nonlinear Dynamics*. Edited by Antonio Ferriz-Mas and Manuel Núñez. London: Taylor and Francis Group, 345
- Berger, M. A. & Field, G. B. 1984, *Journal of Fluid Mechanics* (ISSN 0022-1120), 147, 133
- Biskamp, D. & Müller, W.-C. 1999, *Physical Review Letters*, 83, 2195
- Blackman, E. G. & Hubbard, A. 2014, *Monthly Notices of the Royal Astronomical Society*, 442, 1040
- Brandenburg, A. & Subramanian, K. 2005, *Physics Reports*, 417, 1
- Candelaresi, S. 2012, Ph.D. Thesis, 126
- Christensson, M., Hindmarsh, M., & Brandenburg, A. 2005, *Astronomische Nachrichten*, 326, 393
- Del Sordo, F., Candelaresi, S., & Brandenburg, A. 2010, *Physical Review E*, 81, 036401
- Démoulin, P. 2007, *Advances in Space Research*, 39, 1674
- Démoulin, P. & Berger, M. A. 2003, *Solar Physics*, 215, 203
- Démoulin, P. & Pariat, E. 2009, *Advances in Space Research*, 43, 1013
- DeVore, C. R. 1991, *Journal of Computational Physics* (ISSN 0021-9991), 92, 142

DeVore, C. R. 2000, *The Astrophysical Journal*, 539, 944

Elsasser, W. M. 1956, *Reviews of Modern Physics*, 28, 135

Finn, J. H. & Antonsen, T. M. J. 1985, *Comments on Plasma Physics and Controlled Fusion*, 9, 111

Frisch, U., Pouquet, A., Léorat, J., & Mazure, A. 1975, *Journal of Fluid Mechanics*, 68, 769

Gray, T., Lukin, V. S., Brown, M. R., & Cothran, C. D. 2010, *Physics of Plasmas*, 17, 102106

Heidbrink, W. W. & Dang, T. H. 2000, *Plasma Physics and Controlled Fusion*, 42, L31

Heyvaerts, J. & Priest, E. R. 1984, *Astronomy and Astrophysics*, 137, 63

Ji, H. 1999, *Magnetic Helicity in Space and Laboratory Plasmas*. (1999), 111, 167

Ji, H., Prager, S. C., & Sarff, J. S. 1995, *Physical Review Letters*, 74, 2945

Karpen, J. T., Antiochos, S. K., & DeVore, C. R. 2012, *The Astrophysical Journal*, 760, 81

Knizhnik, K. J., Antiochos, S. K., & DeVore, C. R. 2015, *Joint American Astronomical Society/American Geophysical Union Triennial Earth-Sun Summit*, 1, 11305

Kusano, K., Maeshiro, T., Yokoyama, T., & Sakurai, T. 2004, *The Astrophysical Journal*, 610, 537

Linton, M. G. & Antiochos, S. K. 2002, *The Astrophysical Journal*, 581, 703

Linton, M. G. & Antiochos, S. K. 2005, *The Astrophysical Journal*, 625, 506

Linton, M. G., Dahlburg, R. B., & Antiochos, S. K. 2001, *The Astrophysical Journal*, 553, 905

Liu, Y., Hoeksema, J. T., Bobra, M., et al. 2014a, *The Astrophysical Journal*, 785, 13

Liu, Y., Hoeksema, J. T., & Sun, X. 2014b, *The Astrophysical Journal Letters*, 783, L1

Liu, Y., Zhao, J., & Schuck, P. W. 2013, *Solar Physics*, 287, 279

Low, B. C. 1996, *Solar Physics*, 167, 217

MacNeice, P., Olson, K. M., Mobarry, C., de Fainchtein, R., & Packer, C. 2000, *Computer Physics Communications*, 126, 330

Mandrini, C. H., Pohjolainen, S., Dasso, S., et al. 2005, *Astronomy and Astrophysics*, 434, 725

Mininni, P. D. 2007, *Physical Review E*, 76, 26316

Moffatt, H. K. 1969, *Journal of Fluid Mechanics*, 35, 117

Nakwacki, M. S., Dasso, S., Démoulin, P., Mandrini, C. H., & Gulisano, A. M. 2011, *Astronomy and Astrophysics*, 535, A52

Pariat, E., Antiochos, S. K., & DeVore, C. R. 2009, *The Astrophysical Journal*, 691, 61

Pariat, E., Antiochos, S. K., & DeVore, C. R. 2010, *The Astrophysical Journal*, 714, 1762

Pariat, E., Dalmasse, K., DeVore, C. R., Antiochos, S. K., & Karpen, J. T. 2015, *Astronomy and Astrophysics*, 573, A130

Pariat, E., Démoulin, P., & Berger, M. A. 2005, *Astronomy and Astrophysics*, 439, 1191

Patsourakos, S., Pariat, E., Vourlidis, A., Antiochos, S. K., & Wuelser, J. P. 2008, *The Astrophysical Journal*, 680, L73

Pouquet, A., Frisch, U., & Léorat, J. 1976, *Journal of Fluid Mechanics*, 77, 321

Prager, S. C. 1999, *Magnetic Helicity in Space and Laboratory Plasmas*. (1999), 111, 55

Rachmeler, L. A., Pariat, E., DeForest, C. E., Antiochos, S. K., & Török, T. 2010, *The Astrophysical Journal*, 715, 1556

Rudenko, G. V. & Myshyakov, I. I. 2011, *Solar Physics*, 270, 165

Russell, A. J. B., Yeates, A. R., Hornig, G., & Wilmot-Smith, A. L. 2015, *Physics of Plasmas*, 22, 032106

Rust, D. M. 1994, *Geophysical Research Letters*, 21, 241

Stallard, B. W., Hooper, E. B., Woodruff, S., et al. 2003, *Physics of Plasmas*, 10, 2912

Taylor, J. B. 1974, *Physical Review Letters*, 33, 1139

Taylor, J. B. 1986, *Reviews of Modern Physics*, 58, 741

Thalmann, J. K., Inhester, B., & Wiegmann, T. 2011, *Solar Physics*, 272, 243

Valori, G., Démoulin, P., & Pariat, E. 2012, *Solar Physics*, 278, 347

Valori, G., Démoulin, P., Pariat, E., & Masson, S. 2013, *Astronomy and Astrophysics*, 553, 38

Woltjer, L. 1958, in *Proceedings of the National Academy of Sciences of the United States of America (National Acad Sciences)*, 489–491

Yamada, M. 1999, *Magnetic Helicity in Space and Laboratory Plasmas*. (1999), 111, 129

Yang, S., Büchner, J., Santos, J. C., & Zhang, H. Q. 2013, *Solar Physics*, 283, 369

Zhang, Y., Kitai, R., & Takizawa, K. 2012, *The Astrophysical Journal*, 751, 85

Zhao, L., DeVore, C. R., Antiochos, S. K., & Zurbuchen, T. H. 2015, *The Astrophysical Journal*, 805, 61

Appendix A: Helicity test without the condition:

$$\mathbf{A}(z_{\text{top}}) = \mathbf{A}_p(z_{\text{top}})$$

In the numerical implementation of both the practical DeVore method and the DeVore-Coulomb methods, we impose the condition of Eq. (46), of having the same distribution of \mathbf{A} and \mathbf{A}_p at the top boundary, in addition of the DeVore gauge (Eq. 45) which is valid in the whole volume. Because of this condition, \mathbf{A} and \mathbf{A}_p are both different when computed with one method or the other. The gauge of \mathbf{A} and the gauge of \mathbf{A}_p are linked by Eq. (46). It induces that F_{AAp} is null at the top boundary, and de facto reduces its intensity in the whole domain.

It is however possible to compute and estimate the helicity without enforcing Eq. (46). We can follow the helicity evolution by mixing the vector potential computed with both method. Since gauge invariance of Eq. (10) does not require the use of the same gauge for \mathbf{A} and \mathbf{A}_p , we can use the \mathbf{A} computed with the practical DeVore method with \mathbf{A}_p derived with the DeVore-Coulomb methods. \mathbf{A}_p is thus still satisfying both the DeVore and the Coulomb gauge. As \mathbf{A} , which is only satisfying the DeVore gauge condition, has been computed independently of \mathbf{A}_p , there is no boundary surface along which they share any common distribution. In this appendix we refer to this derivation as the "general DeVore-Coulomb" case.

Fig. A.1, right panel, presents the different terms entering in the decompositions Eqs. (11,15) of the relative helicity. As in Sect. 5.1, the gauge invariance of H_V in this computation relatively to the other methods is ensured with a high precision ($<0.3\%$). As with the others methods, one remarks that H_j and H_{pj} remains constant while H_m , H_{mix} and H_p are different in the general DeVore-Coulomb case. This further confirms the gauge dependance properties of each decomposition.

When looking at the time integrated helicity fluxes (Fig. A.1, right panel), we find again that $H_{\partial V, \text{tot}}$ is following tightly the variation of helicity H_V . As for the two other methods, the helicity dissipation is also very small and with a precision similar to the practical DeVore case (Sect. 4.3). What significantly differs from the two other methods is the repartition of the helicity flux F_{tot} between the different terms which composes it.

Since \mathbf{A}_p is respecting the Coulomb condition the term $dH/dt_{p, \text{var}}$ is null to the numerical precision. Since F_ϕ only involves quantities based on the derivation of the potential field, $H_{\partial V, \phi}$ are equal for both the DeVore-Coulomb and the general DeVore-Coulomb cases. On the other hand, as F_{Bn} , F_{Vn} are only involving \mathbf{A} , $H_{\partial V, Vn}$ and $H_{\partial V, Bn}$ in the general DeVore-Coulomb are equal with their respective estimations in the practical DeVore case.

It is therefore F_{AAp} with concentrate the helicity flux contribution that enable F_{tot} to be quasi gauge-invariant for the three derivations (Fig. A.1, right panel). While $H_{\partial V, \text{AAp}}$ was negligible in both the practical DeVore and the DeVore-Coulomb cases, we observe that this term is now a major contributor of the helicity fluxes. This is not surprising since $H_{\partial V, \text{AAp}}$ results from the existence of large differences between the distribution of \mathbf{A} and \mathbf{A}_p on the boundaries. Both the computations in the practical DeVore and the DeVore-Coulomb methods were enforcing Eq. (46), which induces a very weak value of $H_{\partial V, \text{AAp}}$. We observe that dropping the condition (46) creates a strong $H_{\partial V, \text{AAp}}$.

This test further demonstrates that the choice of the gauge strongly influences the distribution of the helicity fluxes composing the total helicity flux F_{tot} . Only the total flux F_{tot} is a quasi gauge-invariant. None of the terms which are composing the helicity flux F_{tot} shall a priori be neglected. Depending on the gauge, each term can carry a significant contribution. In a

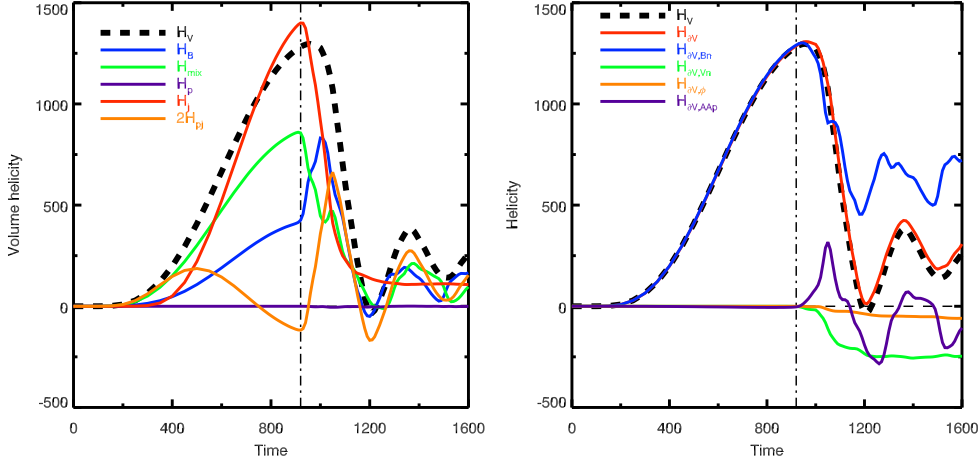


Fig. A.1. **Left panel:** Relative magnetic helicity (H_V , black dashed line) and its decomposition in the general DeVore-Coulomb method. The plotted quantities are the same as in Figures 2, bottom and 7, right panels. **Right panel:** H_V and $H_{\partial V, \#}$ evolution in the system computed with the general DeVore-Coulomb method. The plotted quantities are the same as in Fig. 9.

numerical estimation, it is thus highly advisable to compute all the terms which are forming the helicity flux density (Eq. 23). Explicitly computing each terms allows us to verify that the constraints set on the used gauges are effectively enforced numerically.